



## Dark matter line emission constraints from NuSTAR observations of the Bullet Cluster

Riemer-Sørensen, S.; Wik, D.; Madejski, G.; Molendi, S.; Gastaldello, F.; Harrison, F. A.; Craig, W. W.; Hailey, C. J.; Boggs, S. E.; Christensen, Finn Erland

Total number of authors:  
13

Published in:  
Astrophysical Journal

Link to article, DOI:  
[10.1088/0004-637X/810/1/48](https://doi.org/10.1088/0004-637X/810/1/48)

Publication date:  
2015

Document Version  
Publisher's PDF, also known as Version of record

[Link back to DTU Orbit](#)

### Citation (APA):

Riemer-Sørensen, S., Wik, D., Madejski, G., Molendi, S., Gastaldello, F., Harrison, F. A., Craig, W. W., Hailey, C. J., Boggs, S. E., Christensen, F. E., Stern, D., Zhang, W. W., & Hornstrup, A. (2015). Dark matter line emission constraints from *NuSTAR* observations of the Bullet Cluster. *Astrophysical Journal*, 810, Article 48. <https://doi.org/10.1088/0004-637X/810/1/48>

---

### General rights

Copyright and moral rights for the publications made accessible in the public portal are retained by the authors and/or other copyright owners and it is a condition of accessing publications that users recognise and abide by the legal requirements associated with these rights.

- Users may download and print one copy of any publication from the public portal for the purpose of private study or research.
- You may not further distribute the material or use it for any profit-making activity or commercial gain
- You may freely distribute the URL identifying the publication in the public portal

If you believe that this document breaches copyright please contact us providing details, and we will remove access to the work immediately and investigate your claim.

## DARK MATTER LINE EMISSION CONSTRAINTS FROM *NuSTAR* OBSERVATIONS OF THE BULLET CLUSTER

S. RIEMER-SØRENSEN<sup>1</sup>, D. WIK<sup>2</sup>, G. MADEJSKI<sup>3</sup>, S. MOLENDI<sup>4</sup>, F. GASTALDELLO<sup>4</sup>, F. A. HARRISON<sup>5</sup>, W. W. CRAIG<sup>6,7</sup>, C. J. HAILEY<sup>8</sup>,  
S. E. BOGGS<sup>9</sup>, F. E. CHRISTENSEN<sup>10</sup>, D. STERN<sup>11</sup>, W. W. ZHANG<sup>2</sup>, AND A. HORNSTRUP<sup>10</sup>

<sup>1</sup> Institute of Theoretical Astrophysics, University of Oslo, PO 1029 Blindern, NO-0315 Oslo, Norway; [signe.riemer-sorensen@astro.uio.no](mailto:signe.riemer-sorensen@astro.uio.no)

<sup>2</sup> Astrophysics Science Division, NASA/Goddard Space Flight Center, Greenbelt, MD 20771, USA

<sup>3</sup> Kavli Institute for Particle Astrophysics and Cosmology, SLAC National Accelerator Laboratory, Menlo Park, CA 94025, USA

<sup>4</sup> INFN, IASF Milano, via E. Bassini 15 I-20133 Milano, Italy

<sup>5</sup> Cahill Center for Astronomy and Astrophysics, California Institute of Technology, Pasadena, CA 91125, USA

<sup>6</sup> Space Sciences Laboratory, University of California, Berkeley, CA 94720, USA

<sup>7</sup> Lawrence Livermore National Laboratory, Livermore, CA 94550, USA

<sup>8</sup> Columbia Astrophysics Laboratory, Columbia University, New York, NY 10027, USA

<sup>9</sup> Space Sciences Laboratory, University of California, Berkeley, CA 94720, USA

<sup>10</sup> DTU Space, National Space Institute, Technical University of Denmark, Elektrovej 327, DK-2800 Lyngby, Denmark

<sup>11</sup> Jet Propulsion Laboratory, California Institute of Technology, Pasadena, CA 91109, USA

*Received 2015 July 6; accepted 2015 July 28; published 2015 August 27*

### ABSTRACT

Some dark matter candidates, e.g., sterile neutrinos, provide observable signatures in the form of mono-energetic line emission. We present the first search for dark matter line emission in the 3–80 keV range in a pointed observation of the Bullet Cluster with *NuSTAR*. We do not detect any significant line emission and instead we derive upper limits (95% CL) on the flux, and interpret these constraints in the context of sterile neutrinos and more generic dark matter candidates. *NuSTAR* does not have the sensitivity to constrain the recently claimed line detection at 3.5 keV, but improves on the constraints for energies of 10–25 keV.

*Key words:* dark matter – line: identification – X-rays: galaxies: clusters

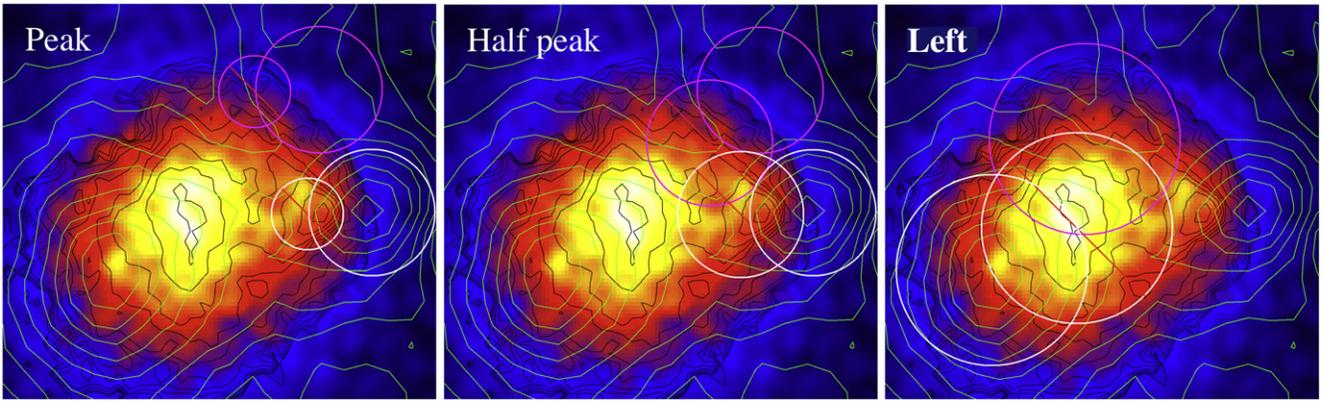
### 1. INTRODUCTION

Dark matter searches are a key pursuit of both astrophysics and particle physics. The scenario where dark matter is in the form of particles that provide gravity, but otherwise interact very weakly with ordinary matter or photons, is the most compelling (e.g., Taoso et al. 2008). The most promising astrophysical objects for searches for dark matter are clusters of galaxies, along with the Galactic Center and dwarf galaxy satellites to the Milky Way. Here, we consider the top end of the mass scale, galaxy clusters, with total masses—most of it in the form of dark matter—often exceeding  $10^{14} M_{\odot}$  ( $\simeq 2 \times 10^{47}$  g). Most cluster mass estimates are inferred from X-ray observations (Vikhlinin et al. 2009; Mantz et al. 2010). Under the assumption of hydrostatic equilibrium, X-ray data provide a reasonably accurate model of the mass distribution and imply cluster masses that are roughly consistent with masses measured via gravitational lensing, albeit both types of data are available for only a limited number of objects. Their total masses are roughly five times greater than the baryonic masses inferred from X-ray luminosities (von der Linden et al. 2014).

One possible particle candidate for dark matter is the sterile neutrino (described in the reviews Boyarsky et al. 2009; Kusenko 2009; Drewes 2013, and references therein). In the framework of this minimally neutrino extension of the standard model ( $\nu$ MSM), the lightest of the three sterile neutrinos provides the dark matter. The mass is basically unbound from theory, but some astrophysical constraints apply. The mass is firmly bound from below through the phase space density of nearby dwarf galaxies. The Tremaine–Gunn bound (Tremaine & Gunn 1979) gives a model-independent lower mass of roughly 0.4 keV (Boyarsky et al. 2008b). This limit can be increased if the production method is known; e.g., for resonant production the boundary is approximately 1 keV (Boyarsky

et al. 2008b; Destri et al. 2013). An upper limit of a few hundred keV comes from a combination of production mechanisms and line emission searches. The sterile neutrino can, in principle, decay spontaneously via mixing with the standard model neutrinos to two photons of equal energy, with the decay probability given by the mixing angle ( $\sin^2(2\theta)$ ). The resulting signature would be a narrow emission line, corresponding to an energy of  $E_{\gamma} = m_s c^2/2$ , where  $m_s$  is the mass of the sterile neutrino. Both the mass and mixing angle are in principle unknown but tied to each other for a given model (Pal & Wolfenstein 1982). The line width would be determined roughly by the velocity dispersion of dark matter particles, which for clusters is of the order of  $v \sim 1000$  km s<sup>-1</sup>, which is smaller than the instrumental resolution of current observatories.

If we can limit the X-ray flux of a line at a specific energy and know the distance, we right away have a limit on the luminosity of the line from the cluster. Since we know the total mass of the cluster, we know how many sterile neutrinos there must be at a given assumed energy to provide the total mass of the cluster. Since we have a limit on the line luminosity, the ratio of the two is basically the limit on the decay rate. A number of authors have reported searches for such emission lines in the soft X-ray band (Abazajian & Koushiappas 2006; Boyarsky et al. 2006a, 2006b, 2007b, 2008a, 2008b; Riemer-Sørensen et al. 2006; Riemer-Sørensen et al. 2007; Loewenstein et al. 2009; Riemer-Sørensen & Hansen 2009; Loewenstein & Kusenko 2010; Leccardi & Molendi 2012), with a few claims of potential detections that are yet to be confirmed (Leccardi & Molendi 2012; Boyarsky et al. 2014; Bulbul et al. 2014). For an unambiguous detection, the searches must avoid the spectral regions with line emission associated with atomic (or nuclear) transitions from the cluster gas,



**Figure 1.** Image of the *NuSTAR* observations of the Bullet Cluster, with the X-ray contours from *Chandra* overlaid in black (Riemer-Sørensen et al. 2007), and the weak lensing contours in green (Clowe et al. 2006). The white circles illustrate the source regions with the second circle excluded to avoid gas emission. The magenta circles are the background regions chosen to have similar gas emission as the source regions, but much less mass.

corresponding to any elements with an expected appreciable cosmic abundance.

Bulbul et al. (2014) recently reported a possible signature for such a sterile neutrino at  $E_\gamma \simeq 3.5$  keV in stacked spectra of galaxy clusters observed with the *XMM-Newton* satellite; the result was confirmed at lower significance in a couple of individual clusters (Boyarsky et al. 2014), but it remains to be independently confirmed in other types of dark matter objects, or using different instruments such as *Suzaku* (Anderson et al. 2014; Malyshev et al. 2014; Riemer-Sørensen 2014; Sekiya et al. 2015; Tamura et al. 2015). Regardless, since there is no theoretical expectation as to the mass of the dark matter sterile neutrino beyond the broad range described above, one should search for its signature in all accessible X-ray spectral bands.

Here, we report on a search extending the energy range to the hard X-ray band, and thus the mass of the putative sterile neutrino to twice the upper end of the *NuSTAR*’s bandpass,  $m_s = 156$  keV. Prior to the launch of *NuSTAR*, there were no sensitive spectral measurements beyond 10 keV, mainly because sensitive measurements require focusing optics (Harrison et al. 2013). While these energies have been searched for line emission previously using the cosmic background (Boyarsky et al. 2006a, 2008b), this is the first search in a pointed observation with focusing optics. The  $E_\gamma > 10$  keV energy range is particularly interesting for entropy-diluted sterile neutrinos (Asaka et al. 2006; Patwardhan et al. 2015).

Recently the *NuSTAR* team observed and reported the results on one well-studied cluster—the “Bullet Cluster” (Wik et al. 2014). There are many previous observations covering its X-ray and lensing properties (Markevitch et al. 2002, 2004; Clowe et al. 2006; Paraficz et al. 2012). This galaxy cluster, at  $z = 0.296$ , is perhaps best known for the detailed comparison of the distribution of dark matter as inferred from gravitational lensing to the X-ray-emitting gas. The lack of spatial alignment between the two distributions reported in Clowe et al. (2006) on the basis of a weak lensing analysis and in Bradač et al. (2006), using a joint weak and strong lensing analysis, is often considered to be one of the strongest arguments for the existence of dark matter particles. Bradač et al. (2006) estimate the total mass of the cluster to be  $5 \times 10^{14} M_\odot$  within the central 500 kpc. Due to the offset between the mass and X-ray-emitting gas, the Bullet Cluster provides an excellent low-background environment for dark matter searches.

The *Chandra* X-ray data of the Bullet have been searched for isolated X-ray emission lines out to 10 keV (Riemer-Sørensen et al. 2007; Boyarsky et al. 2008c) but here we extend the range to  $\sim 80$  keV by using the *NuSTAR* observations from Wik et al. (2014).

Our approach (described in Section 3.1) is to fit the data with an adequate model describing the emission by the hot gas in the cluster, and then search for the improvement to the fit by adding an isolated emission line of varying energy but with fixed width (determined by the instrumental resolution). The detection (or limit on the flux of the line) provide the measurement (or limit) of the flux and thus luminosity of decay photons, yielding the volume-integrated decay rate of the putative dark matter particle (presented in Section 4).

## 2. OBSERVATIONS

The Bullet Cluster was observed by *NuSTAR* in two epochs (see Wik et al. 2014 for details) for a total combined exposure of 266 ks. To filter the events, standard pipeline processing (HEASOFT v6.13 and NUSTARDAS v1.1.1) was applied along with strict criteria regarding passages through the South Atlantic Anomaly (SAA) and a “tentacle”-like region of higher activity near part of the SAA, i.e., in the call to the general processing routine that creates Level 2 data products, `nupipeline`, the following flags were included: `SAAMODE = STRICT` and `TENTACLE = yes`. No strong fluctuations are present in light curves culled from the cleaned events, suggesting a stable background, so no further time periods were excluded.

From the cleaned event files, spectra and response files were created using `nuproducts`. The call to `nuproducts` included `extended = yes`, most appropriate for extended sources, which weights the response files based on the distribution of events within the extraction region, assuming that to be equivalent to the true extent of the source. We extracted spectra for each of the regions shown in Figure 1 (details in Table 1). The regions were chosen to maximize the amount of dark matter within the field of view while minimizing gas emission. The “Peak” region contains the leading mass peak but excludes the shock front. It is identical to the analyzed region in Boyarsky et al. (2008c) and Riemer-Sørensen et al. (2007), while for the “Half Peak” region, the extracted area is larger to compensate for the lower spatial resolution of *NuSTAR* compared to *Chandra*. The “Left” region

**Table 1**  
Details of the Observed Regions

Region	Peak	Half Peak	Left	Peak+Left
Included center (R.A., decl.) (degrees)	104.56825, -55.941758	104.56825, -55.941758	104.64978, -55.951826	...
Included radius (arcmin)	0.66	0.66	1.00	...
Excluded center (R.A., decl.) (degrees)	104.58827, -55.942086	104.59071, -55.994209	104.62326, -55.944488	...
Excluded radius (arcmin)	0.375	0.66	1.00	...
Mass ( $M_{\odot}$ )	$5.7 \times 10^{13}$	$4.5 \times 10^{13}$	$1.0 \times 10^{14}$	$1.6 \times 10^{14}$
Power law $c$ -parameter/dof	1460.3/1355	1195.6/1105	1967.4/1889	3405.2/3221
Background model $c$ -parameter/dof	1610.0/1646	1411.6/1347	2380.3/2315	3990.3/3933

**Note.** The first section provides the coordinates of the regions illustrated in Figure 1. The second section gives the dark matter mass within each field of view based on weak gravitational lensing. The third section contains the fit statistics for each of the regions for the two methods of background subtraction and modeling. The spectra were binned to a minimum of three counts per bin before the analysis.

contains the trailing mass peak with the exclusion of the main gas emission. For each region we chose an offset background region of identical shape at a location with similar gas emission but less mass as inferred from the lensing map. The resulting spectra for all three regions are shown in Figure 2.

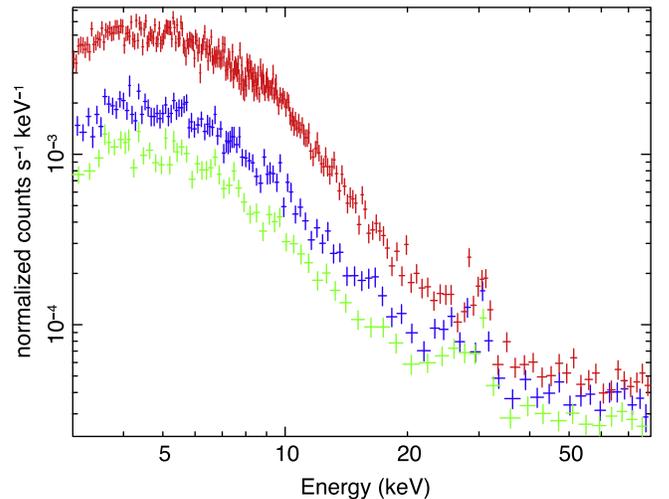
### 3. ANALYSIS

#### 3.1. Spectral Modeling

We fit the spectra of both observations and both detectors simultaneously using the XSPEC spectral fitting package (Arnaud 1996) and explored two different approaches to background modeling and subtraction. The spectra were binned to contain at least three counts per bin, and we assume the background to be Poisson distributed and consequently use Cash-statistics ( $C$ ) to optimize the parameter values (Cash 1979). The Cash-statistics is similar to  $\chi^2$  but generally regarded as more suitable for analyzing spectra with a few counts per channel (Nousek & Shue 1989; Arzner et al. 2007) and less biased than  $\chi^2$  (Leccardi & Molendi 2007). The magnitude of  $C$  depends on the number of bins included in the fit and the values of the data themselves, and consequently one cannot analytically assign a goodness-of-fit measure to a given value of  $C$ . However, for large number counts ( $>5$ ) the  $C$ -distribution is similar to the  $\chi^2$ -statistics and  $\Delta C$  can be used instead of  $\Delta\chi^2$  for model comparison tests.

In the first method of background treatment we simply subtract the spectrum of the dark matter offset region from the spectrum of the source region. Since the spectra are almost identical we can fit any residual with a single power law (the statistics for each region are given in Table 1). Subsequently, we added a Gaussian to represent a single emission line at a fixed energy and flux to the best-fit model above, and searched for the improvement of  $C$  as a function of line energy and intensity, allowing for simultaneous variation of the power-law parameters and the Gaussian normalization. We consider line intensities from 0 to  $10^{-5}$  photons  $\text{cm}^{-2} \text{s}^{-1}$  and line energies of 3–80 keV in steps of  $\Delta E = 0.1$  keV. We assume that the line is narrow compared to the detector resolution, fixing the intrinsic line width at 0.001 keV and noting that as long as the assumed width is less than  $\sim 0.03$  keV, our results do not change.

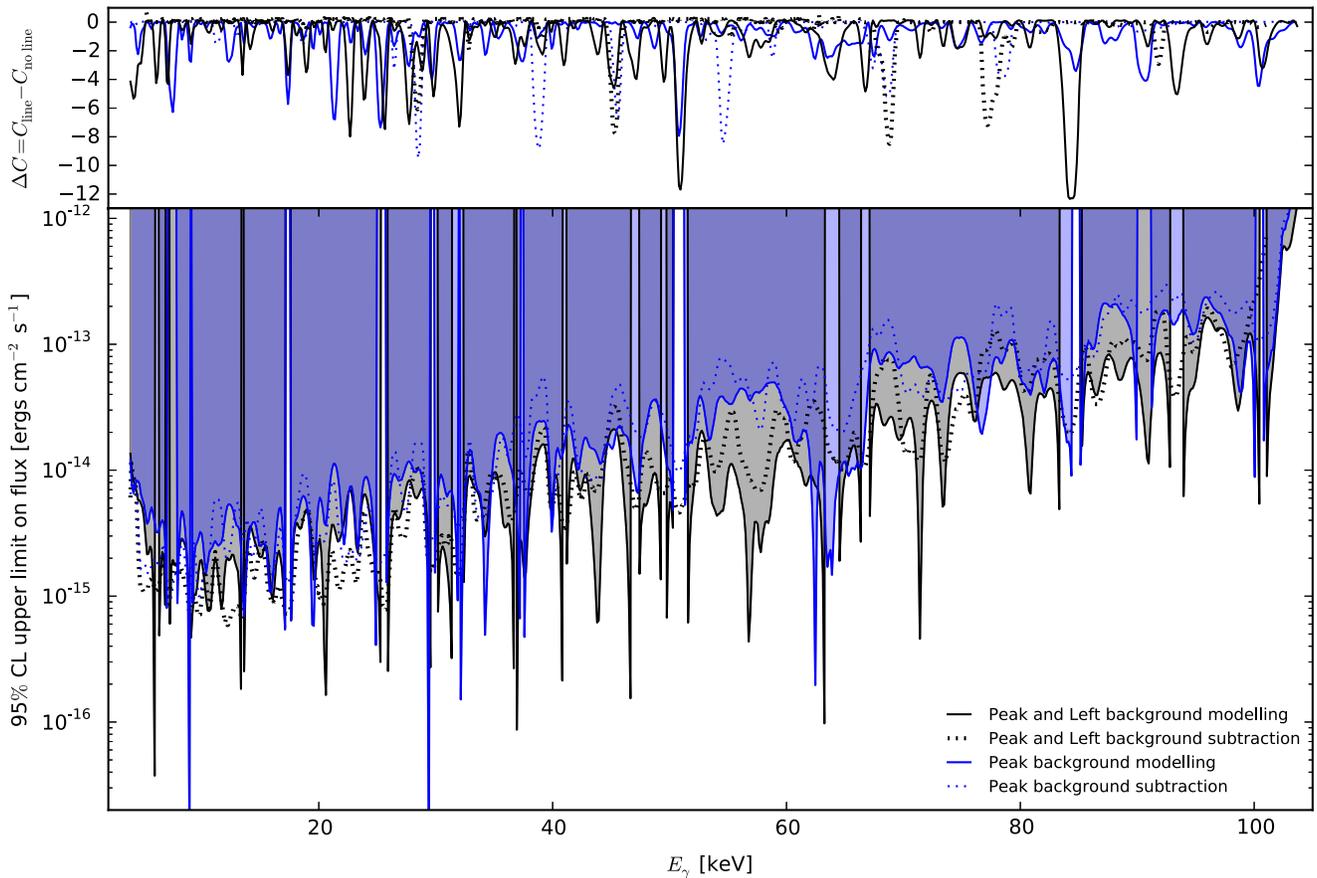
The reduction of the  $C$  parameter by the extra line is shown as a function of line energy in the upper panel of Figure 3. At most energies, the additional Gaussian does not lead to any significant improvement ( $|\Delta C| < 9$  for 1105–3821 degrees of freedom), and instead we constrain the flux by increasing the



**Figure 2.** Spectra of the three different regions: Peak (blue), Half Peak (green), Left (red). The individual exposures have been stacked and rebinned for visualization purposes. The bump around 30 keV comes from instrumental lines and only affects the background modeling method, not the background subtraction.

Gaussian normalization and refitting all other parameters until  $\Delta C = C - C_{\text{base}} = 2.71$ , corresponding to the one-sided 95% confidence level marginalized over the power-law normalizations. These flux levels are shown in Figure 3 for the Peak and Left regions as well as for the combined analysis.

In the second approach we model the background instead of subtracting it. In Wik et al. (2014) the background emission in the Bullet Cluster was thoroughly investigated and we use their results as a baseline model for the background, and check if there is room for any line emission above this model. The model consists of four components: (i) the aperture background (smooth gradient across the detector with a normalization uncertainty of 10%); (ii) a focused cosmic ray background from unresolved sources (smooth with a normalization uncertainty of 10%); (iii) instrumental continuum (we use a 10% normalization uncertainty even though the systematic uncertainty is probably much smaller); and (iv) the thermal solar continuum and instrumental lines from reflections (smooth component with a normalization uncertainty of 10% plus known detector emission lines). Additionally we fit a line-free plasma model (apec; Smith et al. 2001) to account for any gas emission from the cluster. The redshift and abundances of the plasma model are kept fixed at  $z = 0.296$  and  $A = 0.2$ , respectively (consistent



**Figure 3.** Upper panel: the  $C$ -statistics improvement (negative) from adding an extra Gaussian at the corresponding rest frame energy ( $E_\gamma = (1+z)E_{\text{obs}}$ ). Lower panel: the individual derived flux limits (95% CL) for the Peak region (blue) as well as the Peak and Left regions combined (thick black). The solid lines show the results of modeling the background of the nearby dark matter offset region, and the dotted lines show the results of subtracting the nearby dark matter offset region. The results from the individual regions are all consistent with each other and the constraints tighten by combining the regions. The background modeling generally provides slightly stronger constraints, but with gaps where the preferred flux is negative.

with the value found in Wik et al. 2014), while the temperature is required to be the same for both detectors and both observations, but with individual normalizations.

The 3–150 keV energy interval is well fitted by the background model with the statistics given in Table 1. As before, for each energy between 3 and 80 keV in steps of  $\Delta E_\gamma = 0.1$  keV we add a Gaussian and determine the best-fit normalization considering line intensities from  $-10^{-5}$  to  $10^{-5}$  photons  $\text{cm}^{-2} \text{s}^{-1}$ . The line intensities are allowed to be negative to account for overestimation of the background. The normalization is then increased until  $\Delta C^2 = 2.71$ . We derive upper flux limits from the Gaussian alone as well as from the base model plus the Gaussian integrated over the FWHM of the *NuSTAR* spectral resolution crudely approximated by (Harrison et al. 2013)

$$\Delta E_{\text{FWHM}} = 0.01E_\gamma + 0.3 \text{ keV}. \quad (1)$$

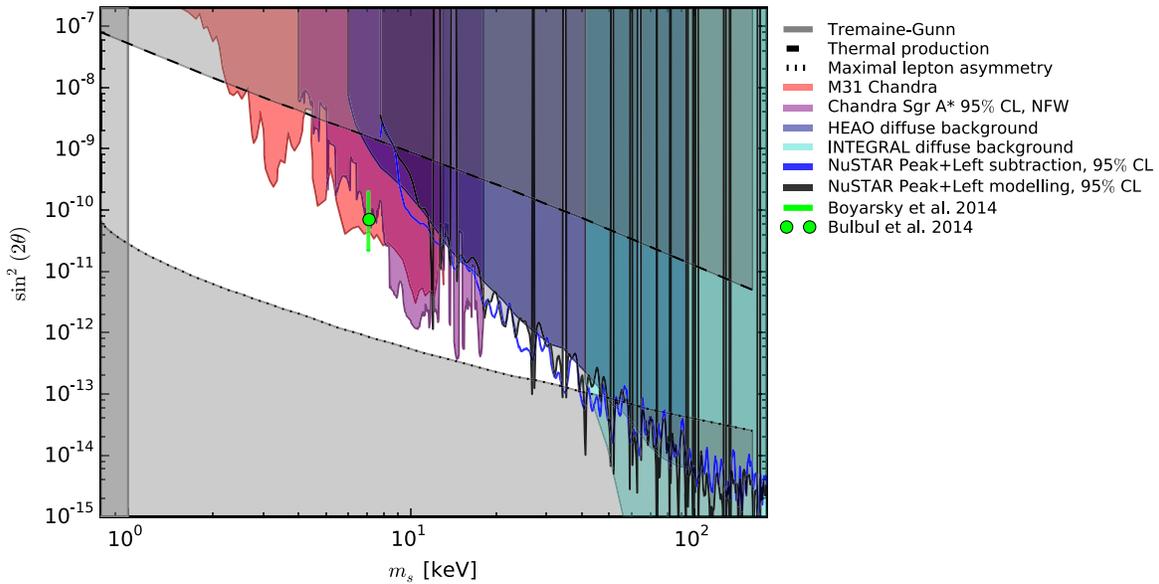
The addition of the Gaussian only improves the model by  $|\Delta C| > 9$  around 50 and 84 keV where  $\Delta C \approx -12$ , reflected in weaker constraints at those energies. While  $|\Delta C| > 9$  would appear to imply a significant detection, we need to take the look-elsewhere effect into account (Gross & Vitells 2010). The energy of the line is unknown and by scanning over energy we perform a number of independent searches that increase the chance of seeing statistical outliers. Consequently the

probability of detection is degraded by the number of independent attempts (given the spectra resolution of *NuSTAR* we search of the order of 150 independent energies).

Line-like features may arise from fluorescent and activation-induced instrumental lines if imperfectly modeled, or due to statistical/systematic fluctuations between target and background regions (see appendix of Wik et al. 2014). These lines are strongest between 20 and 30 keV, and may explain the region of larger  $C$ -values at those energies.

The flux limits from the two approaches for the background subtraction are compared in Figure 3. The solid lines show the results of modeling the background of the nearby dark matter offset region, and the dotted lines show the results of subtracting the nearby dark matter offset region. The results from the individual regions are consistent with each other, and the constraints tighten by combining the regions. The background modeling generally provides slightly stronger constraints, but with gaps where the preferred flux is negative. The results are very similar, and we consider the background treatment to be robust. We also derive flux limits from fitting the two regions simultaneously, taking their mass difference and consequently expected signal strength into account.

While the background modeling provides slightly stronger constraints, there is a risk of spurious line detections from small residuals due to a gain drift between the actual observation and the time of collection of data from which the background



**Figure 4.** Derived constraints on sterile neutrino mixing angle as a function of mass compared to a selection of previous constraints. The upper gray region is excluded by overproduction of dark matter, and the lower one by Big Bang nucleosynthesis (Laine & Shaposhnikov 2008; Canetti et al. 2013). Below  $m_s = 1$  keV the Tremaine–Gunn limit applies (Tremaine & Gunn 1979; Boyarsky et al. 2008b). The green points indicate the potential signal from Bulbul et al. (2014) and Boyarsky et al. (2014) and the colored shaded regions are observation bounds: M31 *Chandra* (Horiuchi et al. 2014), *Chandra* Sgr A\* (Riemer-Sørensen 2014), *HEAO* diffuse background (Boyarsky et al. 2006a), *Fermi* /GBM diffuse background (Ng et al. 2015), and *INTEGRAL* diffuse background (Boyarsky et al. 2008b; Yüksel et al. 2008). The *NuSTAR* Bullet Cluster constraints are similar to previous constraints from the diffuse cosmic background, but provide an important cross-check.

model was constructed. This is unlikely to be the case for direct background subtraction.

### 3.2. Mass within the Field of View

The weak lensing shear maps<sup>12</sup> from Clowe et al. (2006) can be integrated to provide the masses within the three regions for a fiducial cosmology of  $\Omega_m = 0.3$ ,  $\Omega_\Lambda = 0.7$ ,  $H_0 = 70 \text{ km s}^{-1} \text{ Mpc}^{-1}$ . The map contours are shown in Figure 1 and the obtained region masses are given in Table 1.

Paraficz et al. (2012) presented a mass map of the Bullet Cluster based on strong lensing rather than weak lensing. This map provides region masses that are almost twice as big as for the weak lensing map. This indicates an uncertainty on the mass estimates of the order of 50%. In the remaining analysis and plots we conservatively use the values from the map by Clowe et al. (2006) and assume that the entire mass is made up of dark matter since the observed regions have been chosen to minimize the gas presence, so for the Peak region  $M_{\text{gas}} \approx 0.1M_{\text{tot}}$ .

## 4. RESULTS

### 4.1. Sterile Neutrinos

Assuming all of the dark matter to be Majorana-type sterile neutrinos, we can interpret the flux constraints from Figure 3 in terms of the sterile neutrino parameters of mass,  $m_s$ , and mixing angle,  $\sin^2(2\theta)$ , where the latter describes the mixing probability with the lightest of the active neutrinos. The constraints are converted as (Riemer-Sørensen et al. 2006;

Boyarsky et al. 2007b):

$$\sin^2(2\theta) \leq 10^{18} \left( \frac{F_{\text{obs}}}{\text{erg cm}^{-2} \text{ s}^{-1}} \right) \left( \frac{m_s}{\text{keV}} \right) \left[ \frac{(M_{\text{fov}}/M_\odot)}{(D_L/\text{Mpc})^2} \right] \quad (2)$$

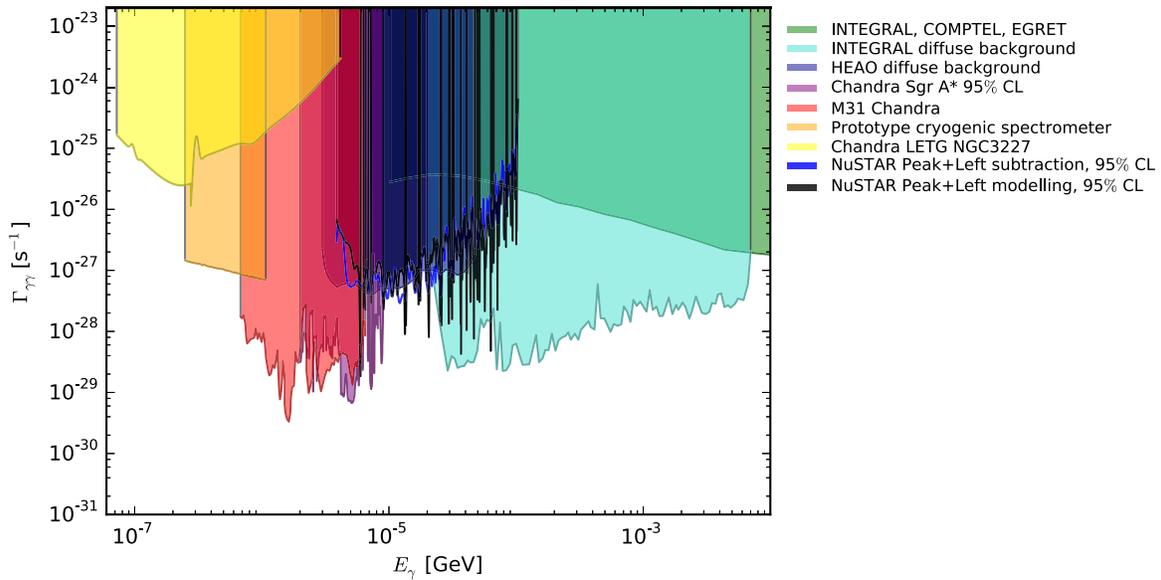
where  $F_{\text{obs}}$  is the observed flux limit,  $M_{\text{fov}}$  is the total dark matter mass within the field of view, and  $D_L$  is the luminosity distance. Equation (2) is only concerned with the decay of the sterile neutrinos and applies regardless of how they were formed. Theoretically, the sterile neutrino can have any mass, but as explained in Section 1 observations restrict the range to approximately 1–100 keV.

As illustrated in Figure 4, very large mixing angles will lead to overproduction of dark matter and are consequently ruled out. Similarly, the resonant production mechanisms require a primordial lepton asymmetry (Boyarsky et al. 2008b), which may affect Big Bang nucleosynthesis and cosmic element abundances. This gives a lower limit on the mixing angle. As mentioned in Section 1, the mass range is limited from below by the Tremaine–Gunn bound (Tremaine & Gunn 1979; Boyarsky et al. 2008b) and from above by line emission searches. The *NuSTAR* Bullet Cluster constraints from the Peak region alone are weaker than previously existing constraints from the diffuse cosmic background, but by combining the Peak and Left regions we get similar constraints. This provides an important independent cross-check of the many assumptions about, e.g., source distribution, that go into the diffuse background constraints.

### 4.2. Generic Dark Matter Constraints

In Figure 5 we present the constraints on generic dark matter decays leading to photon emission for two photons per decay. For one-photon interactions, the constraints are weaker by a

<sup>12</sup> <http://flamingos.astro.ufl.edu/1e0657/index.html>



**Figure 5.** Derived constraints on two-photon decay rates as a function of photon energy compared to previous constraints: *INTEGRAL*, *COMPTEL*, and *EGRET* (Yüksel & Kistler 2008), *INTEGRAL* diffuse background (Boyarsky et al. 2008b), *HEAO* diffuse background (Boyarsky et al. 2006a), *Fermi* /GBM diffuse background (Ng et al. 2015), *Chandra* Sgr A\* (Riemer-Sørensen 2014), M31 *Chandra* (Horiuchi et al. 2014), prototype cryogenic spectrometer (Boyarsky et al. 2007a), and *Chandra* LETG NGC 3227 (Bazzocchi et al. 2008). These constraints apply to all Majorana type dark matter candidates with mono-energetic photon emission in the relevant energy range. The *NuSTAR* Bullet Cluster constraints are similar to previous constraints from the diffuse cosmic background, but provide an important cross-check based on different assumptions.

factor of two. For particles of Majorana-type where the particles are their own anti-particles, the interaction probability doubles and the constraints are strengthened by a factor of two.

Again, the *NuSTAR* constraints do not significantly improve on existing constraints, but as a pointed observation with robust background treatment they provide an important cross-check on previous analysis of diffuse emission observed with *HEAO*, *Fermi*, *INTEGRAL*, *COMPTEL*, and *EGRET* (Boyarsky et al. 2006a, 2008b; Yüksel & Kistler 2008; Yüksel et al. 2008; Ng et al. 2015).

## 5. DISCUSSION

### 5.1. Possible Detections

Unfortunately, the claims of possible line detections (Loewenstein & Kusenko 2010; Boyarsky et al. 2014; Bulbul et al. 2014) all lie below the lower sensitivity cutoff for *NuSTAR* of 3.89 keV introduced by the redshift of the Bullet Cluster.

### 5.2. Possible Improvements

The Milky Way halo has provided strong constraints on line emission in the X-ray range (Riemer-Sørensen et al. 2006; Riemer-Sørensen 2014). The advantage of local constraints is that the dark matter source is nearby and there exists a wealth of observations, but the disadvantage is the number of sources of non-thermal “background” radiation and the uncertainty of the inner mass profile. The background radiation issue can be reduced significantly by point source removal if one has sufficient spatial resolution. This is now becoming possible with *NuSTAR*, and will be investigated further in future work. The profile uncertainty problem can be mitigated somewhat by excluding the center of the halo from the analysis.

## 6. SUMMARY

We have searched *NuSTAR* observations of the Bullet Cluster for exotic line emission over the 3–80 keV interval. No significant line flux was found and we have derived upper limits on the possible line emission flux. While the constraints are similar to previous constraints from the cosmic background emission, this is the first time a search has been performed in this energy interval using a pointed observation. The constraints can be improved by longer observations or by using different targets, e.g., the Milky Way halo (as in Riemer-Sørensen 2014).

This research made use of data from the *NuSTAR* mission, a project led by the California Institute of Technology, managed by the Jet Propulsion Laboratory, and funded by NASA, and it also made use of the *NuSTAR* Data Analysis Software (*NuSTARDAS*) jointly developed by the ASI Science Data Center (ASDC, Italy) and the California Institute of Technology (USA).

*Facility:* *NuSTAR*

## REFERENCES

- Abazajian, K., & Koushiappas, S. M. 2006, *PhRvD*, **74**, 023527  
 Anderson, M. E., Churazov, E., & Bregman, J. N. 2014, arXiv:1408.4115  
 Arnaud, K. 1996, in ASP Conf. Ser. 101, *Astronomical Data Analysis Software and Systems V*, ed. G. H. Jacoby & J. Barnes (San Francisco, CA: ASP), **17**  
 Arzner, K., Güdel, M., Briggs, K., et al. 2007, *A & A*, **468**, 501  
 Asaka, T., Shaposhnikov, M., & Kusenko, A. 2006, *PhLB*, **638**, 401  
 Bazzocchi, F., Lattanzi, M., Riemer-Sørensen, S., & Valle, J. W. F. 2008, *JCAP*, **8**, 13  
 Boyarsky, A., den Herder, J.-W., Neronov, A., & Ruchayskiy, O. 2007a, *Aph*, **28**, 303  
 Boyarsky, A., Iakubovskiy, D., Ruchayskiy, O., & Savchenko, V. 2008a, *MNRAS*, **387**, 1361  
 Boyarsky, A., Malyshev, D., Neronov, A., & Ruchayskiy, O. 2008b, *MNRAS*, **387**, 1345

- Boyarsky, A., Neronov, A., Ruchayskiy, O., & Shaposhnikov, M. 2006a, *MNRAS*, **370**, 213
- Boyarsky, A., Neronov, A., Ruchayskiy, O., Shaposhnikov, M., & Tkachev, I. 2006b, *PhRvL*, **97**, 261302
- Boyarsky, A., Nevalainen, J., & Ruchayskiy, O. 2007b, *A&A*, **471**, 51
- Boyarsky, A., Ruchayskiy, O., Iakubovskiy, D., & Franse, J. 2014, *PhRvL*, **113**, 251301
- Boyarsky, A., Ruchayskiy, O., & Markevitch, M. 2008c, *ApJ*, **673**, 752
- Boyarsky, A., Ruchayskiy, O., & Shaposhnikov, M. 2009, *ARNPS*, **59**, 191
- Bradač, M., Clowe, D., Gonzalez, A. H., et al. 2006, *ApJ*, **652**, 937
- Bulbul, E., Markevitch, M., Foster, A., et al. 2014, *ApJ*, **789**, 13
- Canetti, L., Drewes, M., & Shaposhnikov, M. 2013, *PhRvL*, **110**, 061801
- Cash, W. 1979, *ApJ*, **228**, 939
- Clowe, D., Bradač, M., Gonzalez, A. H., et al. 2006, *ApJL*, **648**, L109
- Destri, C., de Vega, H., & Sanchez, N. 2013, *Aph*, **46**, 14
- Drewes, M. 2013, *IJMP*, **22**, 30019
- Gross, E., & Vitells, O. 2010, *EPJC*, **70**, 525
- Harrison, F. A., Craig, W. W., Christensen, F. E., et al. 2013, *ApJ*, **770**, 103
- Horiuchi, S., Humphrey, P. J., Oñorbe, J., et al. 2014, *PhRvD*, **89**, 025017
- Kusenko, A. 2009, *PhR*, **481**, 1
- Laine, M., & Shaposhnikov, M. 2008, *JCAP*, **6**, 31
- Leccardi, A., & Molendi, S. 2007, *A&A*, **472**, 21
- Loewenstein, M., & Kusenko, A. 2010, *ApJ*, **714**, 652
- Loewenstein, M., & Kusenko, A. 2012, *ApJ*, **751**, 82
- Loewenstein, M., Kusenko, A., & Biermann, P. L. 2009, *ApJ*, **700**, 426
- Malyshev, D., Neronov, A., & Eckert, D. 2014, *PhRvD*, **90**, 103506
- Mantz, A., Allen, S. W., Rapetti, D., & Ebeling, H. 2010, *MNRAS*, **406**, 1759
- Markevitch, M., Gonzalez, A. H., Clowe, D., et al. 2004, *ApJ*, **606**, 819
- Markevitch, M., Gonzalez, A. H., David, L., et al. 2002, *ApJL*, **567**, L27
- Ng, K. C. Y., Horiuchi, S., Gaskins, J. M., Smith, M., & Preece, R. 2015, arXiv:1504.04027
- Nousek, J. A., & Shue, D. R. 1989, *ApJ*, **342**, 1207
- Pal, P. B., & Wolfenstein, L. 1982, *PhRvD*, **25**, 766
- Paraficz, D., Kneib, J.-P., Richard, J., et al. 2012, arXiv:1209.0384
- Patwardhan, A. V., Fuller, G. M., Kishimoto, C. T., & Kusenko, A. 2015, arXiv:1507.01977
- Riemer-Sørensen, S. 2014, arXiv:1405.7943
- Riemer-Sørensen, S., & Hansen, S. H. 2009, *A&A*, **500**, L37
- Riemer-Sørensen, S., Hansen, S. H., & Pedersen, K. 2006, *ApJL*, **644**, L33
- Riemer-Sørensen, S., Pedersen, K., Hansen, S. H., & Dahle, H. 2007, *PhRvD*, **76**, 043524
- Sekiya, N., Yamasaki, N. Y., & Mitsuda, K. 2015, arXiv:1504.02826
- Smith, R. K., Brickhouse, N. S., Liedahl, D. A., & Raymond, J. C. 2001, *ApJL*, **556**, L91
- Tamura, T., Iizuka, R., Maeda, Y., Mitsuda, K., & Yamasaki, N. Y. 2015, *PASJ*, **67**, 23
- Taoso, M., Bertone, G., & Masiero, A. 2008, *JCAP*, **3**, 22
- Tremaine, S., & Gunn, J. E. 1979, *PhRvL*, **42**, 407
- Vikhlinin, A., Kravtsov, A. V., Burenin, R. A., et al. 2009, *ApJ*, **692**, 1060
- von der Linden, A., Allen, M. T., Applegate, D. E., et al. 2014, *MNRAS*, **439**, 2
- Wik, D. R., Hornstrup, A., Molendi, S., et al. 2014, *ApJ*, **792**, 48
- Yüksel, H., Beacom, J. F., & Watson, C. R. 2008, *PhRvL*, **101**, 121301
- Yüksel, H., & Kistler, M. D. 2008, *PhRvD*, **78**, 023502