



Spatially Resolved Water Emission from Gravitationally Lensed Dusty Star-forming Galaxies at $z \approx 3$

Jarugula, Sreevani; Vieira, Joaquin D.; Spilker, Justin S.; Apostolovski, Yordanka; Aravena, Manuel; Bethermin, Matthieu; de Breuck, Carlos; Chen, Chian-Chou; Cunningham, Daniel J. M.; Dong, Chenxing

Total number of authors:
20

Published in:
Astrophysical Journal Supplement Series

Link to article, DOI:
[10.3847/1538-4357/ab290d](https://doi.org/10.3847/1538-4357/ab290d)

Publication date:
2019

Document Version
Publisher's PDF, also known as Version of record

[Link back to DTU Orbit](#)

Citation (APA):

Jarugula, S., Vieira, J. D., Spilker, J. S., Apostolovski, Y., Aravena, M., Bethermin, M., de Breuck, C., Chen, C.-C., Cunningham, D. J. M., Dong, C., Greve, T., Hayward, C. C., Hezaveh, Y., Litke, K. C., Mangian, A. C., Narayanan, D., Phadke, K., Reuter, C. A., Van der Werf, P., & Weiss, A. (2019). Spatially Resolved Water Emission from Gravitationally Lensed Dusty Star-forming Galaxies at $z \approx 3$. *Astrophysical Journal Supplement Series*, 880(2), Article 92. <https://doi.org/10.3847/1538-4357/ab290d>

General rights

Copyright and moral rights for the publications made accessible in the public portal are retained by the authors and/or other copyright owners and it is a condition of accessing publications that users recognise and abide by the legal requirements associated with these rights.

- Users may download and print one copy of any publication from the public portal for the purpose of private study or research.
- You may not further distribute the material or use it for any profit-making activity or commercial gain
- You may freely distribute the URL identifying the publication in the public portal

If you believe that this document breaches copyright please contact us providing details, and we will remove access to the work immediately and investigate your claim.



Spatially Resolved Water Emission from Gravitationally Lensed Dusty Star-forming Galaxies at $z \sim 3$

Sreevani Jarugula¹ , Joaquin D. Vieira^{1,2,3} , Justin S. Spilker⁴ , Yordanka Apostolovski⁵, Manuel Aravena⁵ ,
Matthieu Béthermin⁶ , Carlos de Breuck⁷ , Chian-Chou Chen⁷ , Daniel J. M. Cunningham^{8,9}, Chenxing Dong¹⁰ ,
Thomas Greve^{11,12} , Christopher C. Hayward¹³ , Yashar Hezaveh¹⁴ , Katrina C. Litke¹⁵ , Amelia C Mangian¹,
Desika Narayanan^{10,12,16} , Kedar Phadke¹, Cassie A. Reuter¹, Paul Van der Werf¹⁷ , and Axel Weiss¹⁸

¹ Department of Astronomy, University of Illinois at Urbana-Champaign, 1002 West Green St., Urbana, IL 61801, USA; jarugul2@illinois.edu

² Department of Physics, University of Illinois at Urbana-Champaign, 1110 W Green St Loomis Laboratory, Urbana, IL 61801, USA

³ National Center for Supercomputing Applications, University of Illinois at Urbana-Champaign, 1205 W. Clark St., Urbana, IL 61801, USA

⁴ Department of Astronomy, University of Texas at Austin, 2515 Speedway Stop C1400, Austin, TX 78712, USA

⁵ Núcleo de Astronomía, Facultad de Ingeniería, Universidad Diego Portales, Av. Ejército 441, Santiago, Chile

⁶ Aix Marseille Univ., Centre National de la Recherche Scientifique, Laboratoire d'Astrophysique de Marseille, Marseille, France

⁷ European Southern Observatory, Karl Schwarzschild Straße 2, D-85748 Garching, Germany

⁸ Department of Astronomy and Physics, Saint Mary's University, Halifax, NS, B3H 3C3, Canada

⁹ Department of Physics and Atmospheric Science, Dalhousie University, Halifax, NS, B3H 4R2, Canada

¹⁰ Department of Astronomy, University of Florida, Gainesville, FL 32611, USA

¹¹ Department of Physics and Astronomy, University College London, Gower Street, London WC1E 6BT, UK

¹² Cosmic Dawn Center (DAWN), Niels Bohr Institute, University of Copenhagen, Juliane Maries vej 30, DK-2100 Copenhagen, Denmark; DTU-Space, Technical University of Denmark, DK-2800 Kgs. Lyngby, Denmark

¹³ Center for Computational Astrophysics, Flatiron Institute, 162 Fifth Avenue, New York, NY 10010, USA

¹⁴ Kavli Institute for Particle Astrophysics and Cosmology, Stanford University, Stanford, CA 94305, USA

¹⁵ Steward Observatory, University of Arizona, 933 North Cherry Avenue, Tucson, AZ 85721, USA

¹⁶ University of Florida Informatics Institute, 432 Newell Drive, CISE Bldg E251, Gainesville, FL 32611, USA

¹⁷ Leiden Observatory, Leiden University, PO Box 9513, 2300 RA Leiden, The Netherlands

¹⁸ Max-Planck-Institut für Radioastronomie, Auf dem Hügel 69 D-53121 Bonn, Germany

Received 2019 March 14; revised 2019 May 31; accepted 2019 June 10; published 2019 July 30

Abstract

Water (H_2O), one of the most ubiquitous molecules in the universe, has bright millimeter-wave emission lines that are easily observed at high redshift with the current generation of instruments. The low-excitation transition of H_2O , $p\text{-H}_2\text{O}(2_{0,2} - 1_{1,1})$ ($\nu_{\text{rest}} = 987.927$ GHz), is known to trace the far-infrared (FIR) radiation field independent of the presence of active galactic nuclei (AGNs) over many orders of magnitude in FIR luminosity (L_{FIR}). This indicates that this transition arises mainly due to star formation. In this paper, we present spatially ($\sim 0''.5$ corresponding to ~ 1 kiloparsec) and spectrally resolved (~ 100 kms^{-1}) observations of $p\text{-H}_2\text{O}(2_{0,2} - 1_{1,1})$ in a sample of four strong gravitationally lensed high-redshift galaxies with the Atacama Large Millimeter/submillimeter Array. In addition to increasing the sample of luminous ($> 10^{12} L_{\odot}$) galaxies observed with H_2O , this paper examines the $L_{\text{H}_2\text{O}}/L_{\text{FIR}}$ relation on resolved scales for the first time at high redshift. We find that $L_{\text{H}_2\text{O}}$ is correlated with L_{FIR} on both global and resolved kiloparsec scales within the galaxy in starbursts and AGN with average $L_{\text{H}_2\text{O}}/L_{\text{FIR}} = 2.76^{+2.15}_{-1.21} \times 10^{-5}$. We find that the scatter in the observed $L_{\text{H}_2\text{O}}/L_{\text{FIR}}$ relation does not obviously correlate with the effective temperature of the dust spectral energy distribution or the molecular gas surface density. This is a first step in developing $p\text{-H}_2\text{O}(2_{0,2} - 1_{1,1})$ as a resolved star formation rate calibrator.

Key words: galaxies: high-redshift – galaxies: ISM – galaxies: starburst – ISM: molecules

1. Introduction

Studies of molecules play a prominent role in explaining the physical, chemical, and kinematic properties of the interstellar medium (ISM) in galaxies (Omont 2007; Tielens 2013). One such molecule is H_2O , the third most abundant molecule in the warm dense ISM after H_2 and CO (Neufeld et al. 1995). As an asymmetric rotor with a large electric dipole moment, H_2O has a rich and complex spectrum giving rise to emission and absorption lines mainly in the submillimeter and far-infrared (FIR) regimes of the electromagnetic spectrum. Observations from local galaxies (van der Werf et al. 2010; Weiß et al. 2010; Rangwala et al. 2011; Yang et al. 2013), high-redshift ultraluminous infrared galaxies (ULIRGs; Omont et al. 2013; Yang et al. 2016), and active galactic nuclei (AGNs; van der Werf et al. 2011) have shown H_2O emission to be ubiquitous with intensities as bright as CO lines. Modeling has shown that, in addition to infrared pumping where H_2O is excited by FIR

photons, collisions also contribute to the intensities of low-excitation transitions (e.g., González-Alfonso et al. 2010, 2012). This is best represented in Figure 3 from Liu et al. (2017), which shows prominent H_2O lines in different ISM components. The low-excitation lines become weaker or completely disappear in the warm and hot regions (> 40 K) where infrared pumping dominates over collisions. The higher excitation transitions that require strong far-infrared radiation density are mainly found in the hotter regions (100–200 K) of the galaxy. The cascading emission lines, $p\text{-H}_2\text{O}(2_{0,2} - 1_{1,1})$ ($E_{\text{up}} = 100.8$ K, $\nu_{\text{rest}} = 987.927$ GHz), $p\text{-H}_2\text{O}(2_{1,1} - 2_{0,2})$ ($E_{\text{up}} = 137$ K, $\nu_{\text{rest}} = 752.033$ GHz), and $p\text{-H}_2\text{O}(2_{2,0} - 2_{1,1})$ ($E_{\text{up}} = 196$ K, $\nu_{\text{rest}} = 1228.789$ GHz) are pumped by $101 \mu\text{m}$ photons from the base $1_{1,1}$ level and are primarily excited in the warm regions of the galaxy. The collisional excitation of the low-lying levels ($1_{1,1}$ and $2_{0,2}$) in optically thin or high-density hot regions might also contribute to the emission of the $p\text{-H}_2\text{O}(2_{0,2} - 1_{1,1})$

line. Hence, H_2O transitions probe the infrared radiation field density and physical properties of the ISM such as gas density and kinetic temperature (e.g., Weiß et al. 2010; González-Alfonso et al. 2014; Liu et al. 2017).

Because of water vapor in the Earth’s atmosphere, ground-based observations of H_2O in the local universe are nearly always impossible. The *Herschel Space Observatory* opened the window to multiple H_2O transitions in the local universe (e.g., Weiß et al. 2010) and Yang et al. (2013) demonstrated that the luminosity of submillimeter H_2O lines ($L_{\text{H}_2\text{O}}$) is linearly correlated with the total infrared luminosity (L_{IR} , integrated over 8–1000 μm) over three orders of magnitude in multiple transitions. This suggests that the H_2O transitions, especially $p\text{-H}_2\text{O}(2_{0,2} - 1_{1,1})$, which is not affected by the presence of AGNs (Yang et al. 2013), trace the far-infrared field in star-forming regions. At high redshift, H_2O has been detected using the current generation of ground-based telescopes such as the CSO, PdBI, and ALMA, as the transitions are redshifted into the transparent millimeter atmospheric windows. Strong gravitational lensing, which acts as a cosmic microscope, further boosts the flux from high-redshift sources, making their detections possible. Several detections of H_2O have been reported in the literature from such lensed galaxies (e.g., Bradford et al. 2009; Omont et al. 2011; van der Werf et al. 2011; Combes et al. 2012; Bothwell et al. 2013a; Omont et al. 2013; Weiß et al. 2013; Spilker et al. 2014; Yang et al. 2016).

Multi-wavelength observations ranging from the UV to radio have improved our understanding of interstellar physics and the star formation rate (SFR) calibration. Average scaling relations from single observables are often used to estimate global SFR. Obtaining resolved SFR maps is challenging due to the difficulty of observing individual star-forming regions over multiple wavelengths. The far-infrared luminosity of galaxies (L_{FIR} , integrated over 42.5–122.5 μm) is often used to infer SFR, as it has some advantages over other indicators such as UV luminosity and recombination lines, which are widely discussed in Kennicutt (1998) and Kennicutt & Evans (2012). The UV emission from young stars is a direct tracer of star formation but is highly sensitive to interstellar dust attenuation. Recombination lines such as $\text{H}\alpha$ and FIR cooling lines (e.g., [CII] 158 μm) originate in the ionized regions surrounding stars and are good tracers of star formation. However, these lines are affected either by dust attenuation (e.g., Casey et al. 2017) or the scatter in the estimated SFR is large (e.g., Narayanan & Krumholz 2017; Lagache et al. 2018). In contrast, L_{FIR} is a good tracer of SFR at high optical depth, such as starburst galaxies where most of the UV light is re-emitted as infrared radiation. Although it is widely used as an SFR calibrator in high-redshift starburst galaxies (see review by Casey et al. (2014)), the spectral energy distribution (SED) has to be fully sampled over the SED peak at $\lambda_{\text{rest}} \sim 100 \mu\text{m}$ to estimate L_{FIR} , which is observationally expensive. However, one further caveat is that infrared emission does not necessarily trace only the unobscured star formation. For instance, L_{FIR} may overestimate the SFR in regions where there are other sources of dust heating such as evolved older stars or an obscured AGN (e.g., Kennicutt et al. 2009; Murphy et al. 2011; Hayward et al. 2014). Longer-wavelength spectral features such as $p\text{-H}_2\text{O}(2_{0,2} - 1_{1,1})$, which are very well correlated with L_{FIR} and observable with current generation telescopes, can be used instead of L_{FIR} to estimate SFR (in environments where

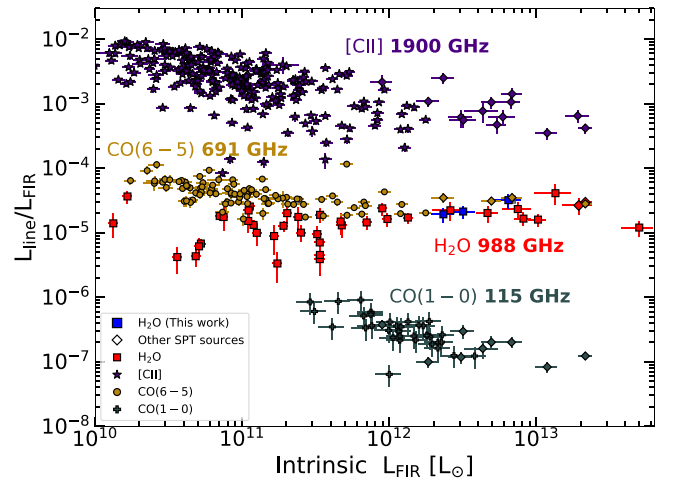


Figure 1. Ratio of line to far-infrared luminosity of $p\text{-H}_2\text{O}(2_{0,2} - 1_{1,1})$ (988 GHz), $\text{CO}(1-0)$ (115 GHz), $\text{CO}(6-5)$ (691 GHz), and [CII] (1900 GHz) as a function of L_{FIR} . The SPT sources are shown as diamonds. The $p\text{-H}_2\text{O}(2_{0,2} - 1_{1,1})$ emission is described in detail in Figure 3. The H_2O emission in the local galaxies is described in Yang et al. (2013) and the emission from high-redshift galaxies is taken from van der Werf et al. (2011), Omont et al. (2013), Yang et al. (2016), Apostolovski et al. (2019), and this paper. $\text{CO}(1-0)$ emission from local ULIRGs is given in Solomon et al. (1997) and the ATCA observations of $\text{CO}(1-0)$ in SPT sources (green diamonds) are described in detail in Aravena et al. (2016). $\text{CO}(6-5)$ line emission from local luminous infrared galaxies (LIRGs) and the SPT sources (golden-yellow diamonds) are from Lu et al. (2017) and Dong et al. (2019), respectively. The [CII] sample of LIRGs is taken from Díaz-Santos et al. (2014) and the [CII] SPT sample represented by purple diamonds is from Gullberg et al. (2015). As seen in the plot, the luminosities of $\text{CO}(1-0)$, $\text{CO}(6-5)$ and [CII] are sublinearly correlated with L_{FIR} , while $p\text{-H}_2\text{O}(2_{0,2} - 1_{1,1})$ is almost linearly correlated with L_{FIR} especially for $L_{\text{FIR}} > 10^{11.5} L_{\odot}$.

the L_{FIR} based calibration holds true). Figure 1 summarizes some of the measurements from the literature and shows that $p\text{-H}_2\text{O}(2_{0,2} - 1_{1,1})$ is almost linearly correlated with L_{FIR} with Pearson’s correlation coefficient of ~ 0.96 . Among the CO transitions, it has been observed that mid to high- J CO transitions (e.g., $\text{CO}(6-5)$ and $\text{CO}(7-6)$) are a good tracer of L_{IR} both in local and high redshift (U)LIRGs (e.g., Lu et al. 2015; Yang et al. 2017). However, sublinear slopes in the $L_{\text{FIR}}\text{-}L_{\text{CO}}$ correlation arising possibly from shocks/turbulence and detached from star formation have also been discussed in high- J CO lines (e.g., Greve et al. 2014, see Section 4.4). While CO is collisionally excited by H_2 molecules, $p\text{-H}_2\text{O}(2_{0,2} - 1_{1,1})$ is excited by FIR photons, which makes H_2O a more direct tracer of star formation. In nearby luminous galaxies, dense gas tracers such as HCN and CS are shown to be tightly correlated with L_{IR} , while HCO^+ has a slightly super-linear correlation (e.g., Gao & Solomon 2004; Zhang et al. 2014). $p\text{-H}_2\text{O}(2_{0,2} - 1_{1,1})$ is a bright emission line (compared to HCN/ HCO^+) that is easily observable both in local and high-redshift galaxies. The linear correlation between $L_{\text{H}_2\text{O}}$ and L_{FIR} from Figure 1 suggests that it is a better tracer of L_{FIR} compared to other commonly observed lines such as $\text{CO}(1-0)$, $\text{CO}(6-5)$, and [CII]. While the correlation is tight on global integrated scales, it is unclear if this correlation breaks down on resolved scales.

In this work, we show that $L_{\text{H}_2\text{O}}$ traces far-infrared radiation not just at the integrated global scale (Omont et al. 2013; Yang et al. 2013, 2016) but also at resolved scales within galaxies at high redshift. The resolution of the observations is $\sim 0''.5$, which corresponds to ~ 1 kiloparsec given the magnification

Table 1
Source Properties and Observations

Source name	IAU name	R.A.	Decl.	z	$\nu_{\text{obs}}^{\text{line}}$	$\nu_{\text{obs}}^{\text{cont}}$	Beam	σ_{cont}	SNR _{line}
		(J2000)	(J2000)		[GHz]	[GHz]	[$''$]	[mJy/ beam]	
SPT0529-54	SPT-S J052903-5436.6	05:29:03.37	−54:36:40.30	3.3689	226.13	224.87	1.10×1.06	0.17	7.5
SPT0532-50	SPT-S J053250-5047.1	05:32:51.27	−50:47:09.50	3.3988	224.59	223.80	0.73×0.56	0.18	24.4
SPT0538-50	SPT-S J053816-5030.8	05:38:16.83	−50:30:52.00	2.7817	260.98	259.32	0.63×0.51	0.22	18.6
Cloverleaf	H 1413+117	14:15:46.24	11:29:43.68	2.5579	277.67	278.83	0.54×0.50	1.07	11.8

Note. The position (R.A., decl.) and redshift (z) of the SPT sources are taken from ALMA 870 μm imaging in Spilker et al. (2016) and Weiß et al. (2013), respectively. The redshift of the Cloverleaf quasar is found in Solomon et al. (2003). $\nu_{\text{obs}}^{\text{line}}$ is the observed frequency of the $p\text{-H}_2\text{O}(2_{0,2} - 1_{1,1})$ transition at a 987.927 GHz rest frequency. $\nu_{\text{obs}}^{\text{cont}}$ is the frequency of the continuum. The beam size and the sensitivity per beam in the continuum map (σ_{cont}) at ν_{obs} are shown in columns 8 and 9. The peak signal-to-noise of the H₂O line (SNR_{line}), with a 50 kms^{-1} channel resolution in all sources except SPT0529-54, with 100 kms^{-1} channel width, is shown in the last column.

and redshift of the sources from Spilker et al. (2014) (the beam resolution and the physical scale for each source are given in Tables 1 and 3, respectively). This physical scale is only an approximation, as we do not perform lens modeling in this analysis and adopt magnification values obtained from 870 μm imaging. We have selected a sample of strong gravitationally lensed dusty star-forming galaxies (DSFGs) discovered in the South Pole Telescope (SPT) survey (Vieira et al. 2010; Carlstrom et al. 2011; Mocanu et al. 2013). DSFGs host intense star formation with $\text{SFR} > 10\text{--}1000 M_{\odot} \text{yr}^{-1}$ (e.g., Casey et al. 2014; Narayanan et al. 2015). These galaxies are bright in submillimeter wavelengths, as the ultraviolet (UV) radiation from young stars is absorbed and re-radiated by the dust in FIR. Long-wavelength dust continuum observations of such galaxies have the advantage of “negative - K correction” (Blain & Longair 1993), where the decrease in flux due to increase in cosmological distance is compensated by the rising flux on the Rayleigh–Jeans side of the SED. Thus, sources of a given luminosity can be detected largely independent of redshift. This, in addition to gravitational lensing and the power of ALMA, provides enough sensitivity and resolution to investigate the correlation between $L_{\text{H}_2\text{O}}$ and L_{FIR} at resolved scales in star-forming galaxies, which we present in this paper.

In Section 2, we summarize the ALMA observations taken over two cycles and the data reduction procedure. In Section 3, we present the results estimating the infrared luminosity and line properties. In Section 4, we analyze the results for the $L_{\text{H}_2\text{O}}\text{--}L_{\text{FIR}}$ correlation and the effect of physical properties on this correlation. We conclude with a summary in Section 5. Hereafter, H₂O refers to $p\text{-H}_2\text{O}(2_{0,2} - 1_{1,1})$ at 987.927 GHz and $L_{\text{H}_2\text{O}}/L_{\text{FIR}}$ refers to $L_{\text{H}_2\text{O}(2_{0,2} - 1_{1,1})}/L_{\text{FIR}}$. We use a Planck 2015 flat ΛCDM cosmology where $h = 0.677$, $\Omega_m = 0.307$, and $\Omega_{\Lambda} = 0.693$ (Planck Collaboration et al. 2016). We estimate the total infrared luminosity (L_{IR}) as flux-integrated from 8 to 1000 μm and total far-infrared luminosity (L_{FIR}) from 42.5 to 122.5 μm in the rest frame (Helou et al. 1985).

2. Observations and Data Analysis

We choose $p\text{-H}_2\text{O}(2_{0,2} - 1_{1,1})$ as it is one of the brightest H₂O transitions and has been observed to be well correlated with L_{IR} (e.g., Yang et al. 2013; Liu et al. 2017). This line also falls in the transparent ALMA Band 6 for the given redshift range of the sources ($z \sim 2.78\text{--}3.37$). We observed the $p\text{-H}_2\text{O}(2_{0,2} - 1_{1,1})$ 987.927 GHz line in SPT0529-54 ($z = 3.369$), SPT0532-50 ($z = 3.399$), and SPT0538-50 ($z = 2.782$) with ALMA. We also

include archival data on the Cloverleaf quasar, a strongly lensed AGN at $z = 2.558$ in this analysis. The source properties are listed in Table 1.

2.1. Sample Selection

The three SPT targets were selected such that they are at a similar redshift and within 10° of each other on the sky. This selection was chosen to observe the same line transition in the three galaxies and to make observations efficient for resolved ALMA Band 9 continuum observations, which were A-rated in Cycle 5, but not yet observed. All three sources have ALMA 870 μm imaging and lens models (Spilker et al. 2016). SPT0538-50 is a possible ongoing major merger as seen from dust continuum models (Bothwell et al. 2013b) and has resolved CO(1–0) and CO(3–2) ATCA observations (Aravena et al. 2013; Spilker et al. 2015). SPT0529-54 and SPT0532-50 have resolved CO(6–5) observations from ALMA (Dong et al. 2019) which we make use of in this work. The Cloverleaf quasar (also known as H1413+117 or QSO J1415+1129) is an extensively studied strongly lensed AGN at high redshift (e.g., Solomon et al. 2003; Weiß et al. 2003) that can be compared against the star-forming galaxies in this sample.

2.2. Data Reduction and Imaging

For SPT0532-50 and SPT0538-50, the observations are taken with ALMA in Cycle 3 (2015.1.01578.S) and Cycle 4 (2016.1.01554.S). SPT0529-54 was observed only in Cycle 4. The Cloverleaf quasar archival data are from ALMA 2012.1.00175.S (PI: Van Der Werf). In both Cycle 3 and Cycle 4, the SPT sources are observed in two basebands each with 2 GHz bandwidth and 240 channels with a channel resolution of 1.875 MHz. The $p\text{-H}_2\text{O}(2_{0,2} - 1_{1,1})$ line and continuum in the Cloverleaf were taken over a continuous bandwidth with two basebands each 2 GHz wide, 128 channels, and a channel resolution of 15.625 MHz. The observations are summarized in Table 2.

All the data are calibrated using the ALMA pipeline for the respective cycles. We inspected the quality of the reduction manually and found no major problems. The data were reduced and imaged using the Common Astronomy Software Application package CASA (McMullin et al. 2007). We imaged the continuum by combining data from all the spectral windows and by excluding the line emission using the task CLEAN. The frequency of the continuum image is given in Table 1. An outer taper of $1''.0$ and $0''.5$ is applied to SPT0529-54 and the

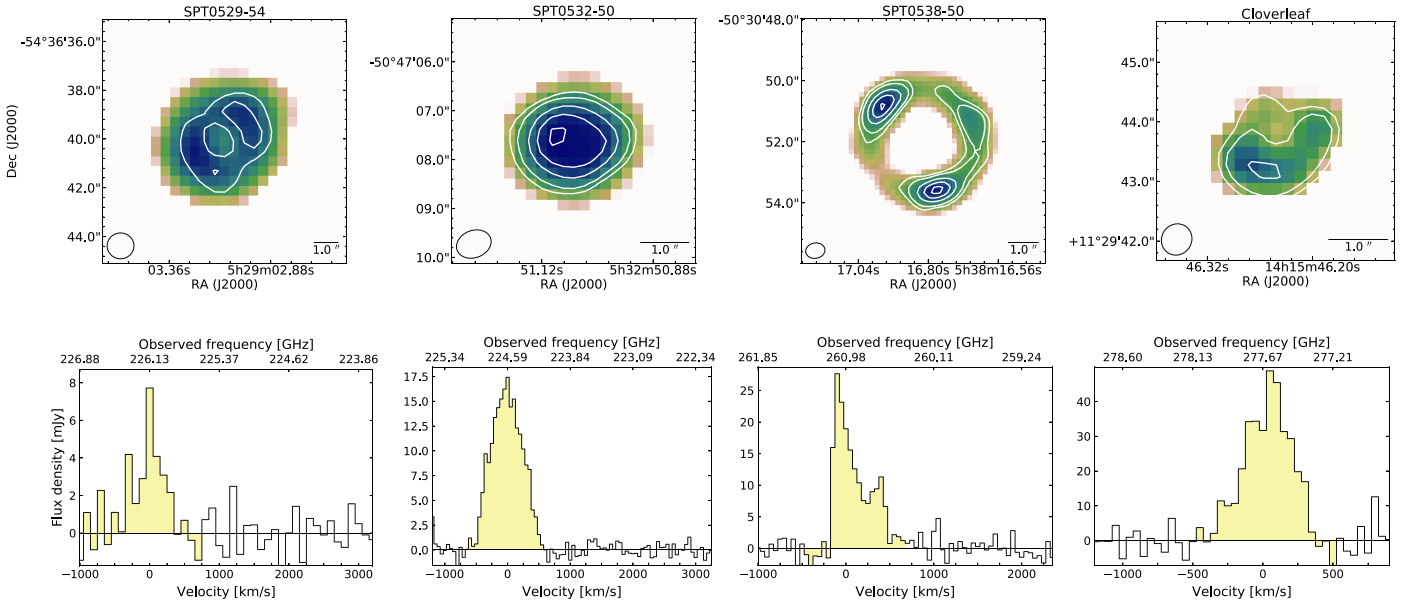


Figure 2. Top: the continuum is shown as a background in log scale with the minimum pixel value as $3\sigma_{cont}$ of the continuum map and the moment 0 contours of the H_2O emission are overlaid in white. The contours are at $[3,5,10,15 \dots] \times \sigma$, where σ is the rms noise in the H_2O moment 0 map. The synthesized beam of the continuum image is shown in the lower left corner and the spatial scale bar of $1.0''$ is shown in the lower right. Bottom: the spatially integrated spectrum of the $p\text{-H}_2\text{O}(2_{0,2} - 1_{1,1})$ transition, with a 50 km s^{-1} spectral resolution in all sources except SPT0529-54 with a 100 km s^{-1} resolution. The colored region shows (line center $-3\sigma_v$) km s^{-1} to (line center $+3\sigma_v$) km s^{-1} (definition given in the text.)

Table 2
ALMA Observations

Cycle	Proposal ID	Source	Date	Time on source (hr)	Antennas	Baseline (m)	PWV (mm)
4	2016.1.01554.S	SPT0529-54	2016 Dec 5	1.01	41	15.1–650.3	1.57–1.84
		SPT0532-50	2016 Dec 3	0.38	40	15.1–704.1	0.53–0.56
		SPT0538-50	2016 Dec 5	0.90	41	15.1–650.3	1.57–1.78
3	2015.1.01578.S	SPT0532-50	2016 Jun 22	0.41	38	15.1–704.1	1.27–1.32
		SPT0538-50	2016 Jun 22	0.74	38	15.1–704.1	1.23–1.28
2	2012.1.00175.S	Cloverleaf	30 June, 2015	0.16	42		

Note. The baseline given in the table is the minimum and maximum. The last column shows the range of Precipitable Water Vapor (PWV) over the course of the observations.

Cloverleaf, respectively, such that the visibilities at shorter baselines are weighted more. This increases the signal-to-noise at the expense of the resolution of the image. A natural weighting is applied to all the sources. All the pixels in the image plane are correlated in the interferometric data. In order to have a minimum number of correlated pixels, we choose to have a few pixels in a beam (~ 3 to 5). The continuum images are shown in Figure 2. We create a mask such that only the pixels with signal-to-noise ≥ 3 in the continuum are selected. We use this mask for all the resolved analysis in this paper.

To get the spectral cubes, we use natural weighting for all the sources and $1''$ and $0.5''$ out tapers to SPT0529-54 and the Cloverleaf, respectively, the same as the continuum map. We use 50 km s^{-1} velocity averaging for SPT0532-50, SPT0538-50 and the Cloverleaf and 100 km s^{-1} in SPT0529-54. For the velocity-integrated intensity map (moment 0), in order to increase the signal-to-noise ratio, we re-image the data to create a single wide channel that contains most of the line flux. The width of this channel is $\sim 2 \times \text{FWHM}$ of the line (see Table 3 for FWHM and integrated line flux values). All the moment 0 maps are imaged

similar to the continuum map and the cube. The moment 0 contours are overlaid on the continuum image in Figure 2.

It has been shown by many studies that the CO gas sizes can be larger than those of infrared emission (e.g., Spilker et al. 2015; Tadaki et al. 2017; Calistro Rivera et al. 2018; Dong et al. 2019). As H_2O and CO are observed to have similar line profiles (e.g., Omont et al. 2013; Yang et al. 2016, and Appendix A3 (C)), it is likely that H_2O and CO are both tracing similar regions. Hence, H_2O could also be more extended than dust. In our analysis, the mask selected from the continuum includes 95%–100% of the total H_2O flux (depending on the source) and any possible additional extended emission would be small and not affect our results. Moreover, selecting a mask based on continuum is less biased because of the high signal-to-noise in every pixel in the continuum, unlike the moment 0 map.

3. Results

3.1. Estimating L_{FIR}

To fit the SED (and estimate the total L_{FIR}) in the SPT sources, we use the unresolved millimeter and submillimeter

Table 3
Observed Continuum and Line Properties

Source	μ	r^a (kpc)	S_{cont} (mJy)	L_{FIR} ($10^{12} L_{\odot}$)	$S_{\text{H}_2\text{O}}^{\text{Peak}}$ (mJy)	$\Delta V_{\text{H}_2\text{O}}$ (km s^{-1})	$I_{\text{H}_2\text{O}}$ (Jy km s^{-1})	$L_{\text{H}_2\text{O}}$ ($10^8 L_{\odot}$)	$L_{\text{H}_2\text{O}}/L_{\text{FIR}}$ (10^{-5})
SPT0529-54	13.2 ± 0.8	2.3	32.9 ± 0.1	30.84 ± 4.96	7.7 ± 1.0	675 ± 235	2.8 ± 0.6	5.94 ± 1.18	1.92 ± 0.49
SPT0532-50	10.0 ± 0.6	1.7	42.3 ± 0.1	68.33 ± 9.81	17.6 ± 0.6	475 ± 44	10.4 ± 0.2	22.25 ± 0.38	3.26 ± 0.47
SPT0538-50	20.1 ± 1.8	1.1	58.3 ± 0.1	63.91 ± 10.55	27.7 ± 1.5	463 ± 144	8.8 ± 1.3	13.59 ± 2.08	2.13 ± 0.48
Cloverleaf	11	0.4	20.1 ± 0.2	70.91 ± 15.47	49.9 ± 3.7	434 ± 156	17.0 ± 2.4	22.63 ± 3.17	3.19 ± 0.83

Note. Magnification (μ) of the SPT sources is taken from ALMA 870 μm lens models from Spilker et al. (2016). Magnification of the Cloverleaf quasar is based on Venturini & Solomon (2003). a) r is the effective resolution in kiloparsecs achieved per beam (beam size is given in Table 1) in the source plane. S_{cont} is the spatially integrated continuum flux ($\geq 3\sigma_{\text{cont}}$) and L_{FIR} is the observed far-infrared luminosity. $S_{\text{H}_2\text{O}}^{\text{Peak}}$ is the peak line flux and $\Delta V_{\text{H}_2\text{O}}$ is the FWHM of the spatially integrated spectrum, which is derived using the non-parametric method of estimating line properties (Section 3.3). $I_{\text{H}_2\text{O}}$ is velocity-integrated line flux under the colored region shown in Figure 2. $L_{\text{H}_2\text{O}}$ is the observed H_2O luminosity. The error on $L_{\text{H}_2\text{O}}/L_{\text{FIR}}$ takes into account both the error on luminosities and magnification.

photometry from ALMA (3 mm), SPT (2.0 and 1.4 mm), LABOCA (870 μm), and *Herschel* (500, 350, 250, 160, and 100 μm) (Vieira et al. 2013; Strandet et al. 2016). For the Cloverleaf, we use the unresolved photometry from Weiß et al. (2003) (references therein). We fit the modified blackbody function given by Equation (1) with the Markov Chain Monte Carlo (MCMC) algorithm using the `emcee` (Foreman-Mackey et al. 2013) package to sample the posterior probability function:

$$S_{\nu} = \frac{\Omega}{(1+z)^3} (B_{\nu}(T_{\text{d}}) - B_{\nu}(T_{\text{CMB}}))(1 - e^{-\tau}), \quad (1)$$

where ν is the rest frequency, Ω is the source solid angle, $B_{\nu}(T_{\text{d}})$ is the Planck function estimated at dust temperature T_{d} , and τ is the optical depth. At long wavelengths, τ is given by $\tau = (\lambda/\lambda_0)^{\beta}$ with λ_0 being the wavelength at which the optical depth is unity (e.g., Draine 2006) and β is the spectral index that determines the slope of the Rayleigh–Jeans tail of the blackbody. This is the same method used previously in, e.g., Greve et al. (2012) and Spilker et al. (2016). It should be noted that this simplistic modified blackbody function applies only with the assumption that the entire source has a single T_{d} and that the source is uniform (e.g., Hayward et al. 2012). Moreover, β , τ , and temperature distribution are degenerate (Papadopoulos et al. 2010).

In this analysis we consider photometry points $\geq 50 \mu\text{m}$ rest frame. The β value converges to ~ 2 for the SPT sample, which has a well-sampled SED. For the SPT sources and the Cloverleaf, we fix β to 2 and let the amplitude ($\Omega/(1+z)^3$), T_{d} , and λ_0 vary. To investigate $L_{\text{H}_2\text{O}}/L_{\text{FIR}}$ at resolved scales, the L_{FIR} for each pixel is obtained by scaling the L_{FIR} of the entire source using the flux contribution of each pixel to the total continuum flux, i.e.,

$$L_{\text{FIR}}^i = \frac{S^i}{S} L_{\text{FIR}}, \quad (2)$$

where i denotes i th pixel and S is the continuum flux obtained by combining all the spectral windows from the observations excluding the line. It should be noted that the resolved L_{FIR} is estimated by scaling the continuum that is not in the FIR regime, i.e., not around the peak of the SED (the frequencies of the continuum are given in Table 1). Hence, the variations in dust temperature, optical depth, etc., across the source are not taken into account in this analysis, and we assume that the galaxies have uniform temperature and opacity distribution.

Improvement on this assumption would require spatially resolved continuum observations that sample the peak of the dust SED at rest frame $\sim 100 \mu\text{m}$ (which as of this date have been approved, but not observed).

3.2. Literature Sample

We draw our sample of sources detected at both low and high redshift, in $p\text{-H}_2\text{O}(2_{0,2} - 1_{1,1})$, from the literature. The local sample is drawn from the *Herschel* Science Archive (Yang et al. 2013) and the AGNs in this sample are identified in Koss et al. (2013). The FIR luminosity in the local galaxies is estimated by fitting 60 and 100 μm photometry (Sanders et al. 2003) with a modified blackbody as discussed in Section 3.1. We fix $\beta = 2.0$ to be consistent with the SPT sources and λ_0 to 100 μm (Draine 2006), as there are only two photometry points available.

The high-redshift ULIRGs include SPT0346-52 (Apostolovski et al. 2019), SPT0125-47 (Appendix A2 B), HFLS3 (Riechers et al. 2013), APM08279+5255 (van der Werf et al. 2011), G12. v2.30, NBV1.78, SDP17b, NAV1.195, and SDP11 (Omont et al. 2013; Yang et al. 2016). L_{FIR} in HFLS3 is estimated using photometry taken from Riechers et al. (2013) and magnification from Cooray et al. (2014). In the quasar APM08279+5255, L_{FIR} is taken from Beelen et al. (2006), Weiß et al. (2007) and is magnification from Riechers et al. (2009). L_{FIR} in all the other sources (except SPT sources) are estimated with the same process used for the local galaxies (fixing β and λ_0) using 250, 350, 500, and 880 μm photometry and magnification from Bussmann et al. (2013). All the high-redshift ULIRG values are given in Table 5.

3.3. Spectral Analysis and $L_{\text{H}_2\text{O}}$

To obtain the spatially integrated flux density in each of the 50 or 100 km s^{-1} wide channels, we apply the continuum mask to each velocity bin of the spectral cube and integrate within the region selected. Selecting a mask based on the continuum flux is less biased because of the high signal-to-noise ratio in every pixel in the continuum, unlike the moment 0 map. This mask includes 95%–100% of the total water emission. The resultant spectra are shown in Figure 2. We use the standard deviation of flux density in line-free channels as the error in each velocity bin. To obtain the line properties, we use the non-parametric estimation of line width described in Bothwell et al. (2013b), which is a more preferred method than fitting a simple Gaussian profile, especially if the line profile is asymmetric. The intensity-weighted second moment of the spectrum is

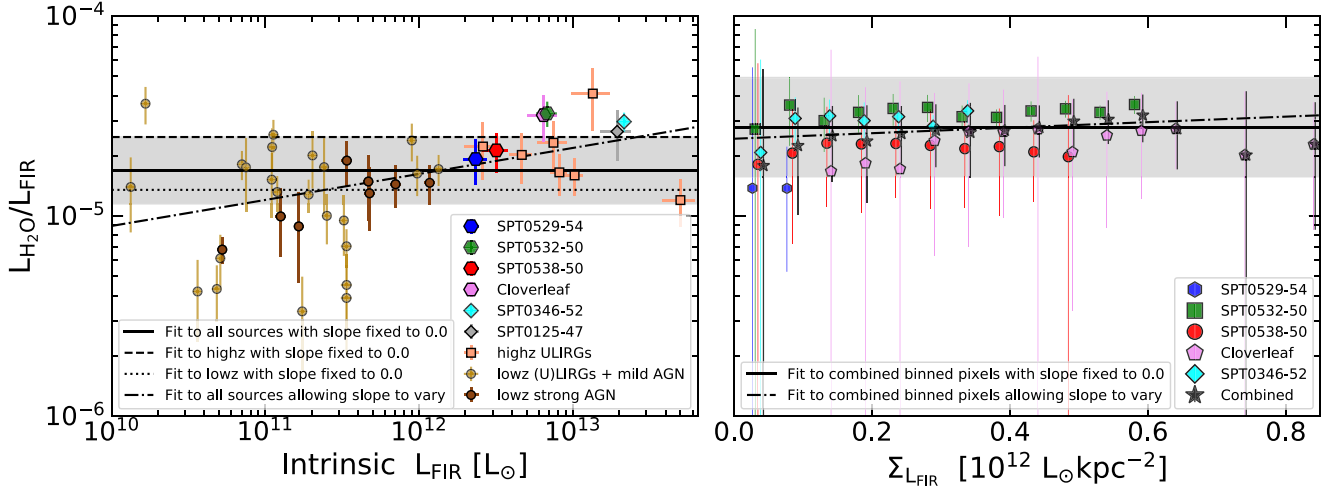


Figure 3. Left (A): global $L_{\text{H}_2\text{O}}/L_{\text{FIR}}$ plotted as a function of spatially integrated intrinsic L_{FIR} (corrected for magnification in high-redshift lensed galaxies). $L_{\text{H}_2\text{O}}$ of the low-redshift LIRGs and ULIRGs (with mild and strong AGNs) are from Yang et al. (2013). $L_{\text{H}_2\text{O}}$ of the high-redshift ULIRGs and AGNs are taken from Omont et al. (2013), Yang et al. (2016), and van der Werf et al. (2011). The three SPT sources, SPT0529-54, SPT0532-50, and SPT0538-50 and the Cloverleaf are presented in this paper. SPT0346-52 is taken from Apostolovski et al. (2019) and SPT0125-47 is presented in Appendix B. The dotted-dashed line is the best fit to all the sources by allowing the slope as a free parameter. The best fit by fixing the slope to zero is shown as a thick black line and the gray region corresponds to the error on the fit. The dashed line is a fit to the high-redshift sources and the dotted line fits the low-redshift galaxies with a fixed slope of zero. Right (B): resolved $L_{\text{H}_2\text{O}}/L_{\text{FIR}}$ plotted as a function of surface brightness in units of $L_{\odot} \text{ kpc}^{-2}$. Each data point is the value of pixels binned within $0.05 \times 10^{12} L_{\odot} \text{ kpc}^{-2}$. The black data points are obtained by combining all the pixels from the five sources and the fit to these points by fixing the slope to zero is shown as a thick black line with the 1σ uncertainty on the fit shown in gray. The dotted line is the best fit by allowing the slope to vary and it is within the gray error region. As shown in the plots, $L_{\text{H}_2\text{O}}$ is strongly correlated with L_{FIR} both at global and resolved scales within the galaxy.

given by

$$\sigma_{\nu} = \frac{\int (\nu - \bar{\nu})^2 S_{\nu} d\nu}{\int S_{\nu} d\nu}, \quad (3)$$

where $\bar{\nu}$ is the intensity-weighted frequency centroid, S_{ν} is the integrated flux at frequency ν , and the FWHM of the line is estimated as $\text{FWHM} \sim 2.35\sigma_{\text{nu}}$. The line properties estimated with this method are listed in Table 3 along with the velocity-integrated line flux ($I_{\text{H}_2\text{O}}$ in Jy km s^{-1}). We obtain the line luminosity from the relation given in Solomon & Vanden Bout (2005):

$$L_{\text{H}_2\text{O}} = (1.04 \times 10^{-3}) I_{\text{H}_2\text{O}} \nu_{\text{rest}} D_L^2 (1+z)^{-1}, \quad (4)$$

where $L_{\text{H}_2\text{O}}$ is the total line luminosity in units of L_{\odot} , ν_{rest} is the rest frequency of the line in GHz (987.927 GHz for the $p\text{-H}_2\text{O}(2_{0,2} - 1_{1,1})$ transition) and D_L is the luminosity distance to the source at a redshift z in Mpc. We estimate $L_{\text{H}_2\text{O}}$ in each pixel using the above equation with $I_{\text{H}_2\text{O}}^i$ (the integrated line flux of the i th pixel taken from the moment 0 image).

4. Analysis and Discussion

4.1. $L_{\text{H}_2\text{O}}\text{-}L_{\text{FIR}}$ Correlation and SFR Calibration

Using the estimated intrinsic luminosities (corrected for magnification) of FIR and H_2O , we plot $L_{\text{H}_2\text{O}}/L_{\text{FIR}}$ as a function of L_{FIR} in Figure 3 in both globally integrated scales and spatially resolved scales (values for high-redshift galaxies are given in Tables 3 and 5 and local galaxies are discussed in Yang et al. 2013). In our analysis, we assume that H_2O and FIR are cospatially lensed, hence the issue of differential lensing where the lensing magnification varies across the source (Blain 1999; Hezaveh et al. 2012) is not significant.

From Figure 3(A), the correlation between global $L_{\text{H}_2\text{O}}$ and L_{FIR} is slightly super-linear and shown as a dotted-dashed black line. A fit (including all sources from the literature in the MCMC) to $\log_{10}(L_{\text{H}_2\text{O}})$ and $\log_{10}(L_{\text{FIR}})$ gives

$$L_{\text{H}_2\text{O}} \propto L_{\text{FIR}}^{(1.17 \pm 0.21)}. \quad (5)$$

This result is similar to the conclusion presented in the literature (Omont et al. 2013; Yang et al. 2016) where $L_{\text{H}_2\text{O}}/L_{\text{FIR}}$ is slightly higher in luminous high-redshift galaxies when compared to less luminous local galaxies. The fit to $L_{\text{H}_2\text{O}}/L_{\text{FIR}}$ with the slope fixed to zero is shown as a thick black line (the error region is shown in gray) gives $L_{\text{H}_2\text{O}}/L_{\text{FIR}} = 1.69_{-0.54}^{+0.79} \times 10^{-5}$. The fit (slope fixed to zero) to high-redshift sources is shown as a dashed line and the fit to the low-redshift galaxies is shown as a dotted line. Both the lines (with their respective errors) are within the error bar of the thick black line, which shows that the increase in $L_{\text{H}_2\text{O}}/L_{\text{FIR}}$ in this sample might not be significant. The almost linear relation between $L_{\text{H}_2\text{O}}$ and L_{FIR} over more than three orders of magnitude supports the previously found conclusions in the literature that $p\text{-H}_2\text{O}(2_{0,2} - 1_{1,1})$ traces L_{IR} both in local galaxies and in high-redshift ULIRGs irrespective of the presence of AGNs.

It should be noted that the low-redshift sample from Yang et al. (2013) is not complete and the super-linear correlation between $L_{\text{H}_2\text{O}}$ and L_{FIR} seen in Figure 3(A) might be real. The deficiency of H_2O in less luminous galaxies ($L_{\text{FIR}} < 10^{11.5} L_{\odot}$) has also been observed in the compilation of fluxes presented by Liu et al. (2017), where a few LIRGs have $L_{\text{H}_2\text{O}}$ similar to ULIRGs while the others either have lower or no H_2O detection. This effect could be arising because of $p\text{-H}_2\text{O}(2_{0,2} - 1_{1,1})$ excitation requirements. H_2O molecules have to be well shielded from UV radiation to avoid dissociation but also have to reside in warm gas (not UV heated) to escape into the gas phase from

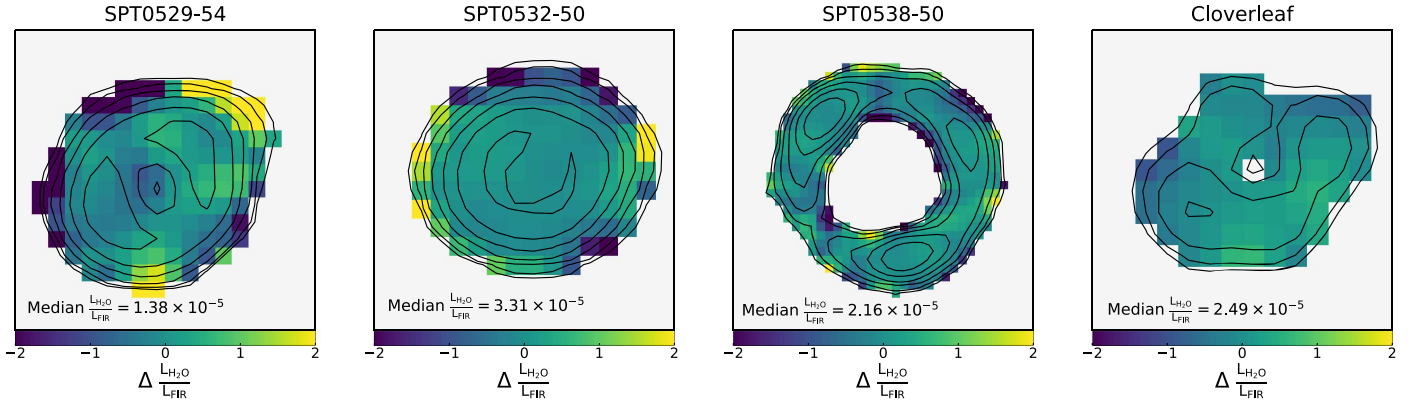


Figure 4. Normalized deviation of $L_{\text{H}_2\text{O}}/L_{\text{FIR}}$ in each pixel from the median value, i.e., $\Delta \frac{L_{\text{H}_2\text{O}}}{L_{\text{FIR}}}$, for each source. The contours in black correspond to continuum emission at $[3, 5, 10, 20, 40, 80 \dots] \times \sigma$, where σ is the rms noise in the continuum map. Values around zero are closer to the median value in that source. The deviation in all the sources is within 10%.

grain mantles. Moreover, $p\text{-H}_2\text{O}(2_{0,2} - 1_{1,1})$ requires dense gas to populate the $1_{1,1}$ level through collisions, which is the base for $101 \mu\text{m}$ excitations. Hence, H_2O couples with far-infrared radiation strongly in such warm, dense, well-shielded gas, which is prevalent in galaxies with $L_{\text{FIR}} > 10^{11.5} L_{\odot}$, but less so in low luminous galaxies. H_2O emission is also enhanced as a result of shocks or intense radiation fields prominent in starbursts. Shocks could increase the abundance of H_2O and strong radiation could lead to an increase in excitation (e.g., Gonzalez et al. 2010; Omont et al. 2011). This might result in the slightly super-linear correlation.

Figure 3(B) shows the correlation between L_{FIR} and $L_{\text{H}_2\text{O}}$ on resolved scales. In each spatial pixel, the surface brightness ($\Sigma_{L_{\text{FIR}}}$) is estimated by dividing the observed L_{FIR} in each pixel by the area of the pixel in the image plane. Because lensing conserves surface brightness, the pixel-by-pixel values of $\Sigma_{L_{\text{FIR}}}$ do not require a lensing correction. $L_{\text{H}_2\text{O}}/L_{\text{FIR}}$ per pixel also does not require a lensing correction, as we assume that the FIR and H_2O emission are cospatially lensed. The data points corresponding to each source are the average value of $L_{\text{H}_2\text{O}}/L_{\text{FIR}}$ in pixels binned within $(0.05 \times 10^{12}) L_{\odot} \text{kpc}^{-2}$ and the error on each data point corresponds to the combined error of the standard deviation of pixel values within that bin and the propagated error due to averaging. This averaging reduces the number of degenerate pixels arising from multiple images of the same region in the source due to gravitational lensing. However, pixel averaging contributes to the error on $\Sigma_{L_{\text{FIR}}}$ and $L_{\text{H}_2\text{O}}/L_{\text{FIR}}$ because of differential lensing within pixels in each bin. These errors are not shown in Figure 3(B), as it is difficult to quantify without a good lens model, but this does not affect the results discussed below. The black points are obtained by combining all the pixels from all the sources. The best fit to the black points by fixing the slope to zero is given by

$$\frac{L_{\text{H}_2\text{O}}}{L_{\text{FIR}}} = 2.76_{-1.21}^{+2.15} \times 10^{-5}. \quad (6)$$

This fit is shown as a thick black line and the gray shaded region in Figure 3(B), which is consistent with the global $L_{\text{H}_2\text{O}}/L_{\text{FIR}}$ (Figure 3(A)). The dotted line in Figure 3(B) is the best fit to the resolved data by allowing the slope to vary. We obtain $L_{\text{H}_2\text{O}}/L_{\text{FIR}} \propto L_{\text{FIR}}^{0.14 \pm 1.44}$, which is consistent within the gray shaded region. Comparing the two fits using a F-test, we conclude that the fit with varying slope does not provide

significantly more information than the fit with a slope fixed to zero. We obtain an F-distribution value of ~ 0.5 with the null hypothesis that the fit with varying slope provides more information than the fit with a slope fixed to zero and we reject this hypothesis if the F-distribution value is > 0.05 .

We observe that $L_{\text{H}_2\text{O}}/L_{\text{FIR}}$ remains linear in SPT0532-50, SPT0538-50, and the Cloverleaf with $\Sigma_{L_{\text{FIR}}}$. The correlation is not obvious in SPT0529-54, as the signal-to-noise ratio of the H_2O emission in this source is lower than those of other sources (peak SNR ~ 7.5 in 100 km s^{-1} channels). This strong correlation between $L_{\text{H}_2\text{O}}$ and $\Sigma_{L_{\text{FIR}}}$ suggests that H_2O is tracing SFR not just at global scales (as discussed in the previous paragraph) but also at resolved scales (~ 1 kiloparsec) within the galaxy. Thus, this result demonstrates that we can use resolved H_2O as a resolved SFR indicator in high-redshift intense star-forming regions. However, high-resolution continuum observations in the FIR regime are required to quantify temperature (and hence L_{IR}) variations within the galaxy.

To understand which regions of the galaxy H_2O best traces L_{FIR} in the image plane at resolved pixel scales, we plot the relative deviation of $L_{\text{H}_2\text{O}}/L_{\text{FIR}}$ in Figure 4. Here, $\Delta \frac{L_{\text{H}_2\text{O}}}{L_{\text{FIR}}}$ corresponds to $[\text{ratio}^i - \text{ratio}_{\text{med}}]/\text{ratio}_{\text{med}}$ where ratio^i is $L_{\text{H}_2\text{O}}/L_{\text{FIR}}$ of i th pixel and $\text{ratio}_{\text{med}}$ is the median value of $L_{\text{H}_2\text{O}}/L_{\text{FIR}}$ in each source. The closer the $\Delta \frac{L_{\text{H}_2\text{O}}}{L_{\text{FIR}}}$ value is to 0, the closer it is to the median value, which implies that H_2O is well correlated with FIR in those pixels. As seen in the images, the correlation is stronger in regions with good signal-to-noise and the ones which deviate the most from $\Delta \frac{L_{\text{H}_2\text{O}}}{L_{\text{FIR}}}$ are at the edges with low signal-to-noise. Moreover, the bright regions are multiple images of the same region in the source (due to gravitational lensing) and as expected, we see that $L_{\text{H}_2\text{O}}/L_{\text{FIR}}$ is the same in these regions. This suggests that water emission faithfully traces the FIR luminosity on resolved scales.

The almost linear correlation between $L_{\text{H}_2\text{O}}$ and L_{FIR} at the resolved scales in the galaxies (Figure 3(B) and Equation (6)) allows us to calibrate the SFR as a function of $L_{\text{H}_2\text{O}}$ for high-redshift intense star-forming regions. The caveat discussed previously in Section 1 that the SFR calibration from L_{FIR} cannot be applied to certain environments, such as regions around AGNs and resolved SFR—depends on variations in dust temperature and opacity. However, we assume a uniform temperature and opacity distribution across the sources, as we

Table 4
Observed Physical Properties in High-redshift ULIRGs

Source	λ_{\max} (μm)	M_{gas} ($10^{10} L_{\odot}$)	A_{eff} (kpc^2)	Σ_{gas} ($10^{10} M_{\odot} \text{kpc}^{-2}$)	References
SPT0529-54	108.57 ± 5.59	4.05 ± 0.90	9.17 ± 3.46	0.44 ± 0.19	Bothwell et al. (2017)
SPT0532-50	92.43 ± 6.83	6.05 ± 1.71	7.46 ± 1.69	0.81 ± 0.29	Bothwell et al. (2017)
SPT0538-50	94.91 ± 6.83	1.7 ± 0.3	15.71 ± 3.05	0.11 ± 0.03	Aravena et al. (2016)
SPT0125-47	84.98 ± 11.17	11.5 ± 1.0	11.93 ± 9.93	0.96 ± 0.81	Aravena et al. (2016)
SPT0346-52	73.80 ± 5.59	8.2 ± 0.6	2.81 ± 0.45	2.92 ± 0.52	Aravena et al. (2016)
Cloverleaf	71.32 ± 6.83

Note. λ_{\max} , the rest-frame wavelength at which dust SED peaks, is estimated from the modified blackbody fit to photometry using the MCMC algorithm. Gas mass (M_{gas}) is taken from the references shown in the last column. The Sérsic area (A_{eff}) is calculated from the best-fit source parameters from lens modeling (Spilker et al. 2016). Σ_{gas} is the gas surface density.

do not have more resolved continuum observations around the peak of the SED. The SFR is generally estimated from the L_{IR} scaling relations discussed in Kennicutt & Evans (2012):

$$\text{SFR} [M_{\odot} \text{yr}^{-1}] = 1.47 \times 10^{-10} L_{\text{IR}} [L_{\odot}]. \quad (7)$$

To convert L_{IR} to L_{FIR} , we use $\langle L_{\text{IR}}/L_{\text{FIR}} \rangle \sim 1.38$ obtained from the SEDs of SPT0529-54, SPT0532-50, SPT0538-50, and the Cloverleaf. This value is similar to $\langle L_{\text{IR}}/L_{\text{FIR}} \rangle$ in the low-redshift galaxies, which is ~ 1.29 (Soifer et al. 1987). Using the L_{IR} to L_{FIR} conversion from the SPT sources and the relation between $L_{\text{H}_2\text{O}}$ and L_{FIR} given by Equation (6), we calibrate SFR using $L_{\text{H}_2\text{O}}$ at resolved galaxy scales as:

$$\text{SFR} [M_{\odot} \text{yr}^{-1}] = 7.35^{+5.74}_{-3.22} \times 10^{-6} L_{\text{H}_2\text{O}} [L_{\odot}]. \quad (8)$$

We have shown that $L_{\text{H}_2\text{O}}$ is well correlated with the FIR continuum at resolved scales. As mentioned previously, this calibration is applicable to high-redshift intense star-forming regions assuming no spatial variations in temperature and optical depth. A similar analysis using resolved continuum observations at the peak of the SED has to be performed to obtain a more accurate SFR calibration from $L_{\text{H}_2\text{O}}$, which could then be used instead of the observationally expensive L_{IR} . We note that resolved FIR continuum observations at the peak of the SED (even at these redshifts) are observationally expensive, while H_2O is a bright line and is easily observable in high-redshift sources with ALMA. In addition to providing an alternative to the expensive FIR continuum observations, H_2O additionally provides kinematics of the star-forming regions.

4.2. Effect of AGNs

Aside from the SPT sources, which are dominated by star formation and show no evidence of an AGN, our sample also includes the Cloverleaf quasar, a well characterized AGN at $z = 2.56$.

The higher transitions of H_2O ($E_{\text{up}} \geq 400 \text{ K}$) are mainly excited by absorption of short-wavelength far-infrared photons ($\leq 50 \mu\text{m}$) emitted by the hot dust surrounding the AGN. Modeling of a lensed quasar, APM 08279+5255 at $z \sim 3.9$ (van der Werf et al. 2011), showed that the higher H_2O transitions are arising from the compact central region with $T_{\text{d}} \sim 200 \text{ K}$. The AGN contributes less to the $J \leq 3$ excitations (mainly excited by $75 \mu\text{m}$ and $101 \mu\text{m}$ photons) in the warm regions but does contribute significantly to the total L_{IR} . This results in APM 08279+5255 lying low on the $L_{\text{H}_2\text{O}}/L_{\text{IR}}$ correlation

(Figure 3(A), Table 5). Similar results are found in Mrk 231 where $p\text{-H}_2\text{O}(4_{2,2} - 4_{1,3})$ is detected (van der Werf et al. 2010) and $L_{\text{H}_2\text{O}}$ in $p\text{-H}_2\text{O}(2_{0,2} - 1_{1,1})$ and $p\text{-H}_2\text{O}(2_{1,1} - 2_{0,2})$ is lower than other ULIRGs without AGNs (Yang et al. 2016).

Previous work (Yang et al. 2013) has shown that although the presence of strong AGN lowers the global $L_{\text{H}_2\text{O}}/L_{\text{IR}}$ ratio, there does not appear to be a significant effect of AGN on H_2O emission lines. This can be seen from Figure 3(A). Our spatially resolved analysis of the Cloverleaf quasar in Figure 3(B) indicates that $L_{\text{H}_2\text{O}}/L_{\text{FIR}}$ remains constant at resolved scales and is similar to ULIRGs even in the presence of an AGN. This suggests that the presence of an AGN has little impact on $p\text{-H}_2\text{O}(2_{0,2} - 1_{1,1})$ excitation—not just at the global scale but also down to kiloparsec scales.

4.3. Correlation of $L_{\text{H}_2\text{O}}/L_{\text{FIR}}$ with Physical Properties

The strength of H_2O emission depends on a number of physical properties such as T_{d} , H_2O column density, and continuum opacity (τ), which are better constrained through modeling of multiple H_2O excitations. As we observed only a single transition, we now investigate to what extent the global variations in $L_{\text{H}_2\text{O}}/L_{\text{FIR}}$ we observe correlate with other properties we constrain such as λ_{\max} (the wavelength in rest frame at which the dust SED peaks) and gas-mass density (Σ_{gas}). We estimate these by using the available photometry and values from the literature (Table 4).

The correlation of $L_{\text{H}_2\text{O}}/L_{\text{FIR}}$ with λ_{\max} can be interpreted as a correlation with the dust temperature or L_{FIR} surface density. Because $p\text{-H}_2\text{O}(2_{0,2} - 1_{1,1})$ is mainly excited by the FIR radiation, we are interested in understanding the correlation with dust temperature. We use λ_{\max} , as it is a more direct observable than dust temperature, which is degenerate with optical thickness (e.g., Papadopoulos et al. 2010). In Figure 5(A), there is no correlation between $L_{\text{H}_2\text{O}}/L_{\text{FIR}}$ and λ_{\max} , consistent with previous results on low-redshift galaxies (Table 2 in Yang et al. 2013), where no relationship is observed between $L_{\text{H}_2\text{O}}/L_{\text{IR}}$ and $S_{60\mu\text{m}}/S_{100\mu\text{m}}$, a dust temperature indicator (these wavelengths are used to fit the SED in our analysis for local sources).

In Figure 5(B), we plot $L_{\text{H}_2\text{O}}/L_{\text{FIR}}$ as a function of Σ_{gas} in an effort to understand whether collisions significantly affect the H_2O excitation. Σ_{gas} is calculated by dividing gas mass by the effective area of the source. The gas-mass values from Bothwell et al. (2017; estimated using [CI] observations) in SPT0529-54 and SPT0532-50 and from Aravena et al. (2016) for SPT0538-50, SPT0125-47, and SPT0346-52 (estimated using CO(1–0) observations) are used. The source properties

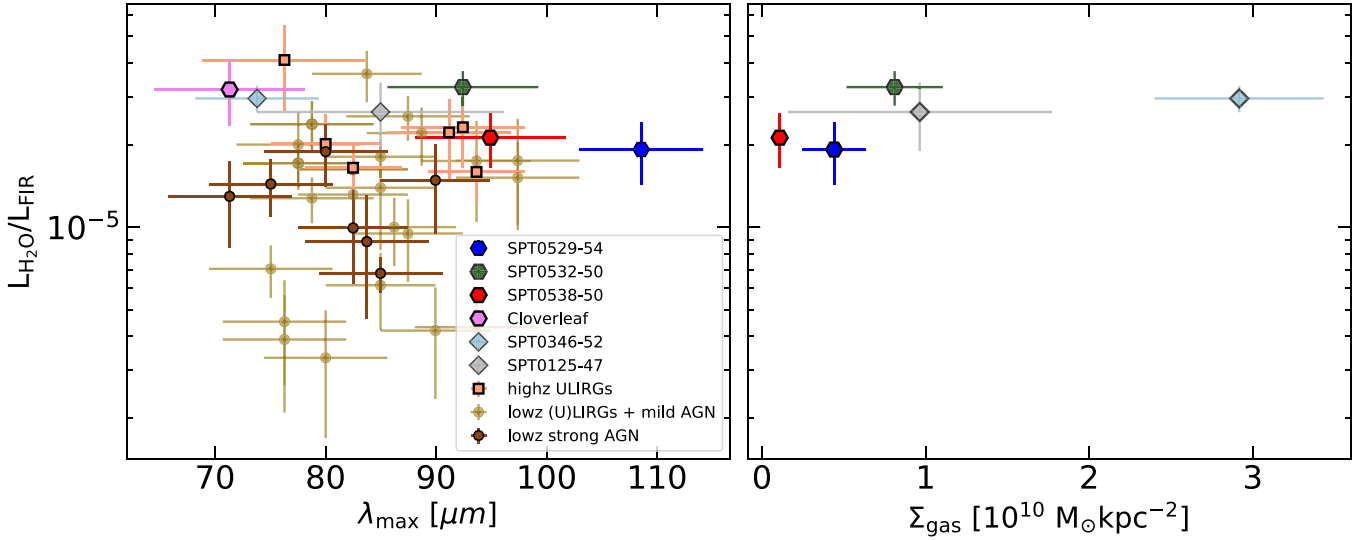


Figure 5. Left (A): correlation between global $L_{\text{H}_2\text{O}}/L_{\text{FIR}}$ and λ_{max} , the rest-frame wavelength at which dust SED peaks. Right (B): global $L_{\text{H}_2\text{O}}/L_{\text{FIR}}$ is plotted as a function of gas surface density, Σ_{gas} in units of $M_{\odot} \text{ kpc}^{-2}$. The gas masses for SPT0529-54 and SPT0532-50 are taken from Bothwell et al. (2017). SPT0538-50 and other two SPT sources (SPT0125-47 and SPT0346-52) are detailed in Aravena et al. (2016). The intrinsic SPT source sizes obtained from lens modeling can be found in Spilker et al. (2016). As shown in the plots, the variation in $L_{\text{H}_2\text{O}}/L_{\text{FIR}}$ is uncorrelated with either the effective temperature of the dust SED or the gas surface density.

are detailed in Spilker et al. (2016), where lens modeling of the $870 \mu\text{m}$ dust continuum is performed by assuming single or multiple Sérsic source profiles. Using these values, the area (A_{eff}) under a Sérsic profile is calculated. This method might overestimate Σ_{gas} , but because the sizes are within a factor of ~ 2 , the overestimated value might only be by a factor of few. Moreover, the CO (gas) sizes can be larger than that of the infrared emission (Spilker et al. 2015; Tadaki et al. 2017; Calistro Rivera et al. 2018; Dong et al. 2019). For a simple calculation, we assume that the dust and gas sizes are similar. All the values are given in Table 4. As seen in Figure 5(B), there is no observed correlation of $L_{\text{H}_2\text{O}}/L_{\text{FIR}}$ with Σ_{gas} .

The dust opacity at $100 \mu\text{m}$, the wavelength at which $p\text{-H}_2\text{O}(2_{0,2} - 1_{1,1})$ is excited, could also affect the intensity of the line and could explain the spread in the $L_{\text{H}_2\text{O}}\text{-}L_{\text{FIR}}$ correlation. The slightly super-linear $L_{\text{H}_2\text{O}}/L_{\text{FIR}}$ correlation where the increase in H_2O line emission is faster than L_{FIR} (e.g., Omont et al. 2013; Yang et al. 2016) could be because of the increase in τ_{100} (dust opacity at $100 \mu\text{m}$) with the increase in L_{FIR} (González-Alfonso et al. 2014), in turn enhancing $L_{\text{H}_2\text{O}}$ because of photon trapping. In a high τ_{100} medium, the $100 \mu\text{m}$ photons are trapped and scattered thereby increasing the local radiation field. This amplifies the $p\text{-H}_2\text{O}(2_{0,2} - 1_{1,1})$ pumping and hence the $p\text{-H}_2\text{O}(2_{0,2} - 1_{1,1})$ line photons. We do a simple estimation of τ_{100} in the three SPT sources using the equation from Yang et al. (2016) where τ_{100} is given by

$$\tau_{100} = \kappa_{100} \frac{M_{\text{dust}}}{2\pi r^2}, \quad (9)$$

κ_{100} is the dust absorption opacity at $100 \mu\text{m}$ and r is the radius of the source at submillimeter wavelengths. We use $\kappa_{\lambda} = 2.92 \times 10^5 (\lambda/\mu\text{m})^{-2} \text{ cm}^2 \text{ g}^{-1}$ (Li & Draine 2001) at rest wavelength λ and dust mass (M_{dust}) given by

$$M_{\text{dust}} = \mu^{-1} \frac{D_L^2 S_{\nu}}{(1+z_s) \kappa_{\lambda} [B_{\nu}(T) - B_{\nu}(T_{\text{CMB}})]}. \quad (10)$$

Here, S_{ν} is the flux density at observed frequency, z_s is the redshift of the source, D_L is the luminosity distance, $B_{\nu}(T)$ is

the Planck function at rest frequency (described in Section 3.1), and μ is the magnification of the source. From the lens model parameters derived in Spilker et al. (2016), we estimate $\tau_{100} \sim 0.34, 1.36$ and 0.46 for SPT0529-54, SPT0532-50, and SPT0538-54, respectively. The higher value of τ_{100} in SPT0532-50 could be enhancing the H_2O luminosity above the average value. However, several physical factors like T_d , opacity, H_2O abundance, etc., can also influence the intensity of the line.

To summarize, the global variations in $L_{\text{H}_2\text{O}}/L_{\text{FIR}}$ are not observed to be affected by the physical properties of the galaxy such as λ_{max} and Σ_{gas} . Large dust opacity at $100 \mu\text{m}$ might enhance $L_{\text{H}_2\text{O}}$ due to photon trapping. Modeling with multiple transitions will give a better understanding of the factors influencing the correlation between L_{FIR} and $L_{\text{H}_2\text{O}}$.

4.4. H_2O and CO

CO(6–5) traces relatively dense gas (with critical density of $\text{H}_2 \sim 10^5 \text{ cm}^{-3}$) in molecular clouds, although not as dense as HCN or HCO^+ (Shirley 2015; Béthermin et al. 2018). The high $-J$ CO lines are therefore found to be correlated with the far-infrared field in these star-forming regions (Figure 1 in Liu et al. 2015). Here, we investigate this correlation in the context of the $L_{\text{H}_2\text{O}}\text{-}L_{\text{FIR}}$ relation. We make use of the spatially and spectrally resolved observations of mid- J CO(6–5) in SPT0529-54, SPT0532-50 (Dong et al. 2019), and supplement these data with observations of the two other SPT sources SPT0346-52 (Apostolovski et al. 2019) and SPT1247-50 (which is not detected in $p\text{-H}_2\text{O}(2_{0,2} - 1_{1,1})$) from Dong et al. (2019). The imaging of the CO data is similar to that described in Section 2.2. The mask used to select the pixels recovers 93%–100% of the CO emission, depending on the source.

Figure 6 shows $L_{\text{CO}(6-5)}/L_{\text{FIR}}$ as a function of L_{FIR} similar to Figure 3. Figure 6(A) contains the global integrated values in local luminous infrared galaxies from Lu et al. (2017) and high-redshift SPT sources. Figure 6(B) shows the resolved correlation between CO(6–5) and L_{FIR} (the binning procedure

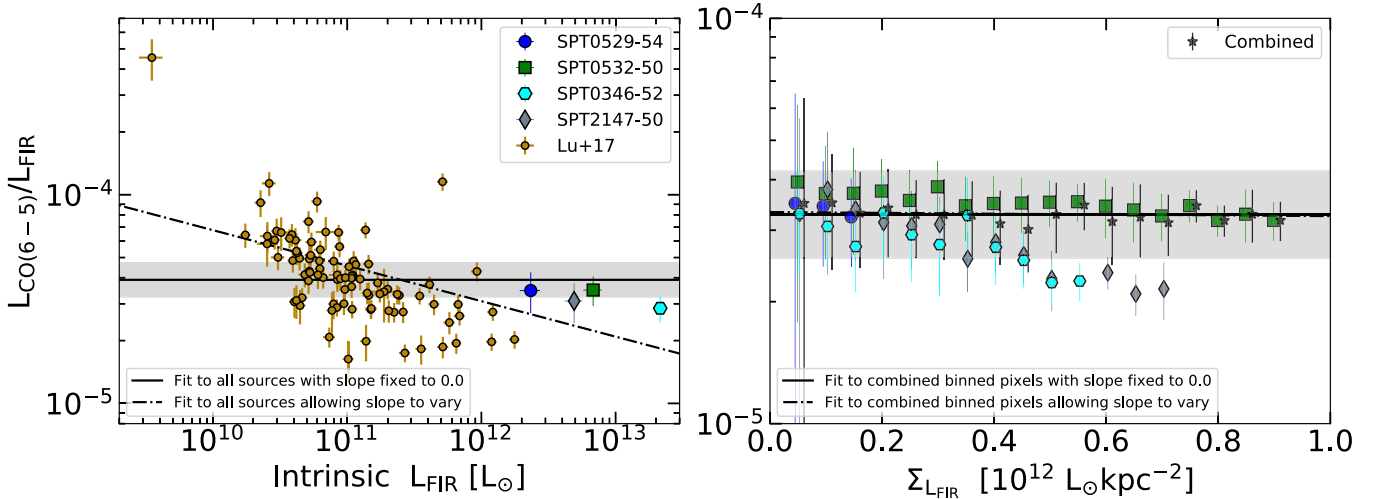


Figure 6. Left (A): global $L_{\text{CO}(6-5)}/L_{\text{FIR}}$ as a function of L_{FIR} . The local LIRGs are shown as yellow data points taken from Lu et al. (2017). The high-redshift ULIRGs are represented by the SPT sources. The thick black line is a fit to all the sources by fixing the slope to zero, with the 1σ error shown as the gray region. The dotted-dashed line is the fit from allowing the slope to vary. From the plot, we see that the correlation is almost linear. Right (B): resolved $L_{\text{CO}(6-5)}/L_{\text{FIR}}$ as a function of surface brightness in units of $L_{\odot} \text{kpc}^{-2}$. Each data point is the value of pixels binned within $0.05 \times 10^{12} L_{\odot} \text{kpc}^{-2}$. The fits are to the combined binned pixels shown in black. The correlation within the sources follow a similar pattern as the global values in the left plot and is also nearly linear. This plot and Figure 3 together suggest that H_2O is as good a tracer of the far-infrared radiation as CO(6–5).

is similar to that described in Section 4.1), where the resolved L_{FIR} is estimated using the continuum around the CO(6–5) line. CO(6–5) is observed to have an almost linear correlation with L_{FIR} both at global and resolved scales, similar to H_2O .

The spectra of CO(6–5) and $p\text{-H}_2\text{O}(2_{0,2} - 1_{1,1})$ in SPT0532-50 and SPT0346-52 (with a good detection of both the lines) show that CO has an FWHM consistent with H_2O within the errors (Figure 8(A)). This may indicate that both the lines are emitted from similar regions in the galaxy (See also Omont et al. 2013; Yang et al. 2016; Liu et al. 2017). It can further be seen from the spatial distribution comparison in the image plane (Figure 8(B)). This agrees with the results in Yang et al. (2019), where they find similar spatial distribution and also similar kinematic structure between CO(6–5) and $p\text{-H}_2\text{O}(2_{1,1} - 2_{0,2})$ in G09v1.97.

While $p\text{-H}_2\text{O}(2_{0,2} - 1_{1,1})$ excitation is due to FIR pumping mechanism and depends mainly on the radiation field density, the CO excitation is due to collisions with the H_2 molecules. Hence, CO intensity increases with increase in the gas density and temperature (e.g., Narayanan & Krumholz 2014). The mid and high- J CO lines ($J = 6-5$ and above) are shown to have increasingly sublinear slopes with L_{FIR} , which suggests that denser and much warmer gas than star-forming regions (possibly due to shocks or turbulence) is needed for CO excitation (Narayanan et al. 2008; Greve et al. 2014). Thus, variations in star formation efficiency, or other physical properties within the molecular gas such as the shape of the density probability distribution function (PDF) and the median density of the gas within and between galaxies might affect $L_{\text{CO}(6-5)}/L_{\text{FIR}}$ more strongly than $L_{\text{H}_2\text{O}}/L_{\text{FIR}}$. Although the mid- J $L_{\text{CO}}/L_{\text{FIR}}$ ratio is not expected to be enhanced in galaxies with supernovae or stellar wind driven shocks, NGC 6240 shows a higher ratio (Lu et al. 2017). This suggests that H_2O is an intrinsically better tracer of the far-infrared field than CO(6–5). To confirm this result, we need a larger sample of sources across a broad range in L_{FIR} to compare H_2O and CO (and other dense gas tracers such as HCN) to determine which one is an empirically better tracer of star formation.

5. Summary and Conclusion

We observed the $p\text{-H}_2\text{O}(2_{0,2} - 1_{1,1})$ 987.927 GHz line in SPT0529-54 ($z = 3.369$), SPT0532-50 ($z = 3.399$), and SPT0538-50 ($z = 2.782$) with ALMA. We also include the Cloverleaf quasar at $z = 2.558$ to compare with the star-forming galaxies. The observational results and conclusions from this analysis are as follows:

1. $L_{\text{H}_2\text{O}}$ is empirically correlated with L_{FIR} over more than three orders of magnitude from low-redshift LIRGs to high-redshift ULIRGs.
2. The relationship between $L_{\text{H}_2\text{O}}$ and L_{FIR} stays linear even at resolved scales within individual galaxies with average $L_{\text{H}_2\text{O}}/L_{\text{FIR}} = 2.76^{+2.15}_{-1.21} \times 10^{-5}$.
3. This linear correlation holds even in the presence of a strong AGN in the Cloverleaf quasar.
4. We present $p\text{-H}_2\text{O}(2_{0,2} - 1_{1,1})$ as a resolved SFR calibrator for high-redshift intense star-forming regions, assuming a single temperature and opacity across the source

$$\text{SFR} [M_{\odot} \text{yr}^{-1}] = 7.35^{+5.74}_{-3.22} \times 10^{-6} L_{\text{H}_2\text{O}} [L_{\odot}].$$

5. There is no observed correlation of $L_{\text{H}_2\text{O}}/L_{\text{FIR}}$ with λ_{max} , the wavelength at which SED peaks or Σ_{gas} , the gas-mass surface density. The dust opacity at $100 \mu\text{m}$ (τ_{100}), on the other hand, may influence $L_{\text{H}_2\text{O}}$ due to photon trapping. However, the current sample is too small to give any definite result.
6. $p\text{-H}_2\text{O}(2_{0,2} - 1_{1,1})$ is intrinsically a better tracer of L_{FIR} than CO(6–5). A larger sample size is needed to confirm this result.

This work shows that $p\text{-H}_2\text{O}(2_{0,2} - 1_{1,1})$ traces L_{FIR} at resolved \sim kiloparsec scales in high-redshift galaxies with intense star-forming regions while assuming a single temperature and dust opacity across the source. In order to validate these assumptions and obtain a more accurate SFR calibration,

Table 5
Observed Properties in High-redshift ULIRGs

Source	z	μ	λ_{max} (μm)	L_{FIR}/μ ($10^{12} L_{\odot}$)	$L_{\text{H}_2\text{O}}/\mu$ ($10^8 L_{\odot}$)	$L_{\text{H}_2\text{O}}/L_{\text{FIR}}$ (10^{-5})	References
SPT0125-47	2.5148	5.467 ± 0.120	84.98 ± 11.17	19.40 ± 4.37	5.12 ± 0.87	2.64 ± 0.74	Appendix A2 B
SPT0346-52	5.6559	5.570 ± 0.117	73.80 ± 5.59	21.50 ± 2.31	6.36 ± 0.24	2.96 ± 0.32	Apostolovski et al. (2019)
G12.v2.30	3.259	9.5 ± 0.6	82.49 ± 4.34	8.16 ± 1.02	1.35 ± 0.27	1.65 ± 0.39	Omont et al. (2013)
NAv1.195	2.951	4.1 ± 0.3	93.67 ± 4.34	10.25 ± 1.38	1.63 ± 0.27	1.59 ± 0.34	Yang et al. (2016)
SDP11	1.786	10.9 ± 1.3	91.18 ± 5.59	2.60 ± 0.63	0.58 ± 0.12	2.22 ± 0.71	Yang et al. (2016)
NBv1.78	3.111	13.0 ± 1.5	80.01 ± 4.97	4.65 ± 0.79	0.94 ± 0.21	2.02 ± 0.57	Omont et al. (2013)
SDP17	2.305	4.9 ± 0.7	92.43 ± 5.59	7.47 ± 1.65	1.73 ± 0.32	2.32 ± 0.67	Omont et al. (2013)
HFLS3	6.337	2.2 ± 0.3	76.28 ± 7.45	13.47 ± 3.74	5.51 ± 1.12	4.09 ± 1.41	Riechers et al. (2013)
APM08279 +5255	3.9	4.0		50.0 ± 12.0	6.0 ± 1.2	1.2 ± 0.3	van der Werf et al. (2011)

Note. For the SPT sources, z and magnification (μ) are given in Spilker et al. (2016) and λ_{max} (the wavelength in rest frame at which the dust SED peaks) and L_{FIR} are estimated by fitting a modified blackbody function to the photometry by fixing $\beta = 2.0$. H_2O observations of SPT0346-52 is discussed in detail in Apostolovski et al. (2019). In HFLS3, photometry is from Riechers et al. (2013) and magnification is from Cooray et al. (2014). The magnification in APM08279+5255 is taken from Riechers et al. (2009). For all other sources, μ and photometry are from Bussmann et al. (2013). L_{FIR} and λ_{max} are estimated by fixing $\beta = 2.0$ and $\lambda_0 = 100 \mu\text{m}$ except in APM08279+5255, where L_{FIR} is taken from (Beelen et al. 2006; Weiß et al. 2007). $L_{\text{H}_2\text{O}}$ is taken from the references given in the last column.

we need resolved continuum observations around the peak of the SED. We also need to perform a similar analysis on less luminous galaxies ($L_{\text{FIR}} < 10^{12} L_{\odot}$) to extend the SFR calibration. Future work will involve detailed lens modeling of the sources with a pixelated lens model (Hezaveh et al. 2016). In the future, it would also be interesting to compare and model multiple resolved H_2O lines with other dense gas tracers.

This paper makes use of the following ALMA data: ADS/JAO. ALMA #2015.1.01578.S, #2016.1.01554.S, #2012.1.00844.S, #2012.1.00175.S, and #2011.0.00957.S. The SPT is supported by the NSF through grant PLR-1248097, with partial support through PHY-1125897, the Kavli Foundation, and the Gordon and Betty Moore Foundation grant GBMF 947. S.J., J.D.V., D.P.M., and K.C.L. acknowledge support from the US NSF under grants AST-1715213 and AST-1716127. S.J. and K.C.L. acknowledge support from the US NSF NRAO under grants SOSPA5-001 and SOSPA4-007, respectively. J.D.V. acknowledges support from an A. P. Sloan Foundation Fellowship. D.N. was supported in part by NSF Award AST-1715206 and *HST* Theory Award 15043.0001. ALMA is a partnership of ESO (representing its member states), NSF (USA) and NINS (Japan), together with NRC (Canada), MOST and ASIAA (Taiwan), and KASI (Republic of Korea), in cooperation with the Republic of Chile. The Joint ALMA Observatory is operated by ESO, AUI/NRAO, and NAOJ. This research has made use of NASA’s Astrophysics Data System.

Appendix A

Appendix: A1

The high redshift ULIRGs sample taken from the literature is shown in Table 5.

Appendix B

Appendix: A2 APEX Observations of H_2O in SPT0125-47

We observed $p\text{-H}_2\text{O}(2_{0,2} - 1_{1,1})$ ($\nu_{\text{rest}} = 987.927$ GHz) line in SPT0125-47 at $z = 2.5148$ using the APEX-2 receiver of the Swedish Heterodyne Facility Instrument (Vassilev et al. 2008) on the Atacama Pathfinder Experiment (APEX). The observations in the shared ESO+Swedish project 092.A-0467 (PI M. Aravena) were done between July and November 2013 in excellent conditions with precipitable water vapor $0.25 < \text{PWV} < 0.5$ mm,

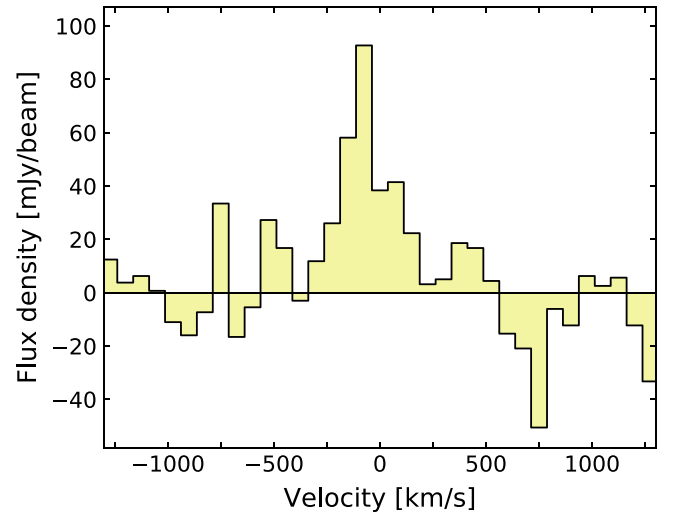


Figure 7. Spatially integrated spectrum of the $p\text{-H}_2\text{O}(2_{0,2} - 1_{1,1})$ transition with a 75 km s^{-1} spectral resolution in SPT0125-47.

and a total on-source integration time of 3 hours. We reduced the data using the standard procedures in the IRAM CLASS software. The line is clearly detected (Figure 7), with a line flux of $21.8 \pm 3.7 \text{ Jy km s}^{-1}$ and a line width of $\sim 117 \text{ km s}^{-1}$. Note that the source is unresolved in the 280 GHz APEX beam of $22''.3$.

Appendix C

Appendix: A3 Spatial distribution of CO(6–5) and H_2O in SPT0532-50 and SPT0346-52

Figure 8 shows the spatial distribution of CO(6–5) and $p\text{-H}_2\text{O}(2_{0,2} - 1_{1,1})$ in two high-redshift ULIRGs, SPT0532-50 and SPT0346-52. From the spatial distribution comparison we can see that both the lines are tracing similar regions in the velocity space. Although the source is gravitationally lensed, the similar distribution in the image plane might indicate that they are tracing the same regions of the galaxy in the source plane.

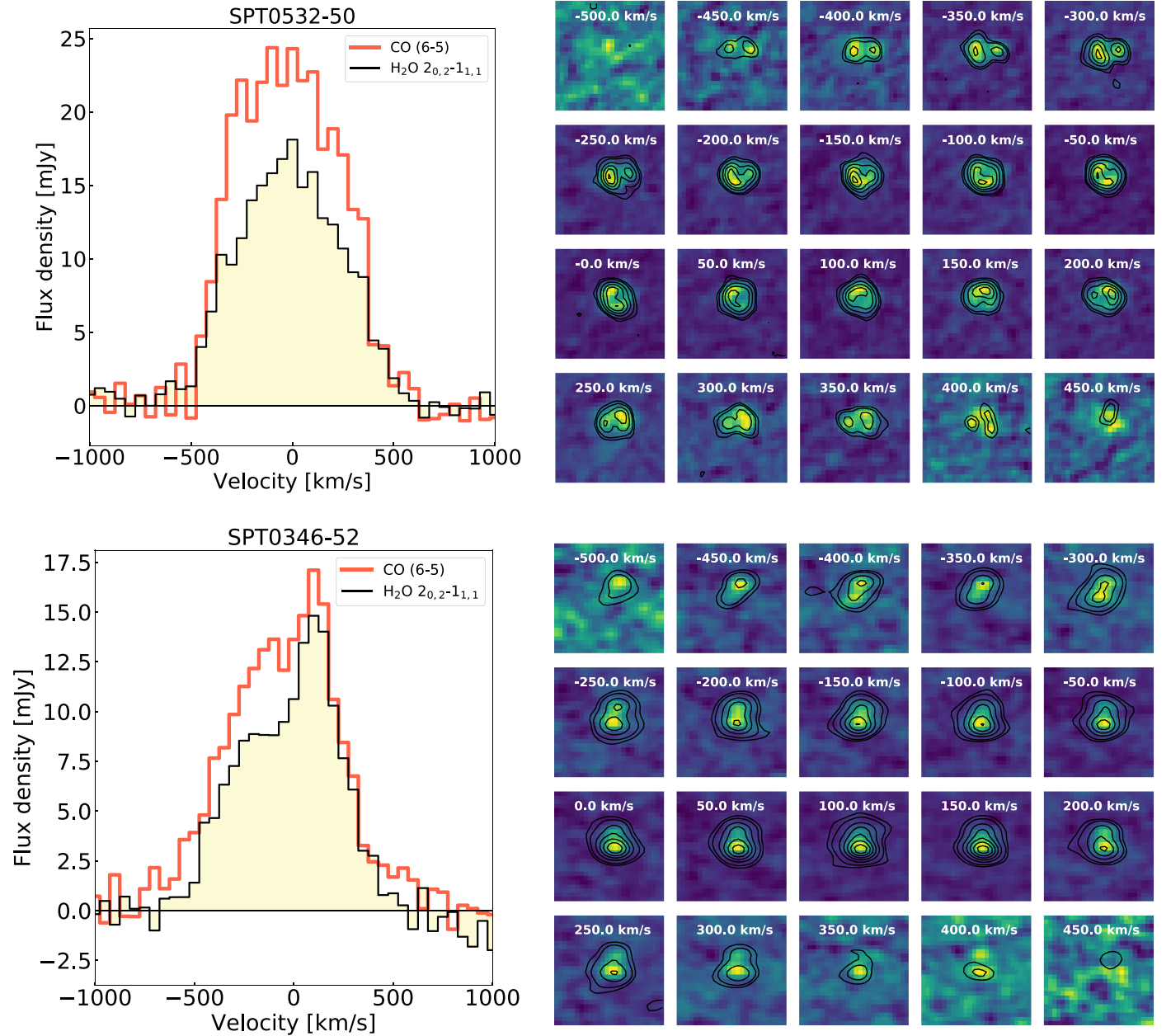


Figure 8. Left (A): spectra of CO(6–5) and p -H₂O(2_{0,2} – 1_{1,1}) in SPT0532-50 and SPT0346-52 integrated over 50 km⁻¹ channels. Both CO and H₂O have similar FWHMs. Right (B): the channel map of p -H₂O(2_{0,2} – 1_{1,1}) in the background and CO(6–5) contours overlaid in black in both the sources. The contours are at [3, 5, 10, 15 ...] × σ , where σ is the rms noise in the CO cube. CO(6–5) and H₂O may be tracing similar warm dense regions of the galaxy.

ORCID iDs

Sreevani Jarugula <https://orcid.org/0000-0002-5386-7076>
 Joaquin D. Vieira <https://orcid.org/0000-0001-7192-3871>
 Justin S. Spilker <https://orcid.org/0000-0003-3256-5615>
 Manuel Aravena <https://orcid.org/0000-0002-6290-3198>
 Matthieu Béthermin <https://orcid.org/0000-0002-3915-2015>
 Carlos de Breuck <https://orcid.org/0000-0002-6637-3315>
 Chian-Chou Chen <https://orcid.org/0000-0002-3805-0789>
 Chenxing Dong <https://orcid.org/0000-0002-5823-0349>
 Thomas Greve <https://orcid.org/0000-0002-2554-1837>
 Christopher C. Hayward <https://orcid.org/0000-0003-4073-3236>
 Yashar Hezaveh <https://orcid.org/0000-0002-8669-5733>
 Katrina C. Litke <https://orcid.org/0000-0002-4208-3532>

Desika Narayanan <https://orcid.org/0000-0002-7064-4309>
 Paul Van der Werf <https://orcid.org/0000-0001-5434-5942>
 Axel Weiss <https://orcid.org/0000-0003-4678-3939>

References

Apostolovski, Y., Aravena, M., Anguita, T., et al. 2019, arXiv:1905.12738
 Aravena, M., Murphy, E. J., Aguirre, J. E., et al. 2013, *MNRAS*, 433, 498
 Aravena, M., Spilker, J. S., Béthermin, M., et al. 2016, *MNRAS*, 457, 4406
 Beelen, A., Cox, P., Benford, D. J., et al. 2006, *ApJ*, 642, 694
 Béthermin, M., Greve, T. R., De Breuck, C., et al. 2018, *A&A*, 620, A115
 Blain, A. W. 1999, *MNRAS*, 304, 669
 Blain, A. W., & Longair, M. S. 1993, *MNRAS*, 264, 509
 Bothwell, M. S., Aguirre, J. E., Aravena, M., et al. 2017, *MNRAS*, 466, 2825
 Bothwell, M. S., Aguirre, J. E., Chapman, S. C., et al. 2013a, *ApJ*, 779, 67
 Bothwell, M. S., Smail, I., Chapman, S. C., et al. 2013b, *MNRAS*, 429, 3047
 Bradford, C. M., Aguirre, J. E., Aikin, R., et al. 2009, *ApJ*, 705, 112

- Bussmann, R. S., Pérez-Fournon, I., Amber, S., et al. 2013, *ApJ*, 779, 25
- Calistro Rivera, G., Hodge, J. A., Smail, I., et al. 2018, *ApJ*, 863, 56
- Carlstrom, J. E., Ade, P. A. R., Aird, K. A., et al. 2011, *PASP*, 123, 568
- Casey, C. M., Cooray, A., Killi, M., et al. 2017, *ApJ*, 840, 101
- Casey, C. M., Narayanan, D., & Cooray, A. 2014, *PhR*, 541, 45
- Combes, F., Rex, M., Rawle, T. D., et al. 2012, *A&A*, 538, L4
- Cooray, A., Calanog, J., Wardlow, J. L., et al. 2014, *ApJ*, 790, 40
- Díaz-Santos, T., Armus, L., Charmandaris, V., et al. 2014, *ApJL*, 788, L17
- Dong, C., Spilker, J. S., Gonzalez, A. H., et al. 2019, *ApJ*, 873, 50
- Draine, B. T. 2006, *ApJ*, 636, 1114
- Foreman-Mackey, D., Hogg, D. W., Lang, D., & Goodman, J. 2013, *PASP*, 125, 306
- Gao, Y., & Solomon, P. M. 2004, *ApJ*, 606, 271
- Gonzalez, A. H., Papovich, C., Bradač, M., & Jones, C. 2010, *ApJ*, 720, 245
- González-Alfonso, E., Fischer, J., Aalto, S., & Falstad, N. 2014, *A&A*, 567, A91
- González-Alfonso, E., Fischer, J., Graciá-Carpio, J., et al. 2012, *A&A*, 541, A4
- González-Alfonso, E., Fischer, J., Isaak, K., et al. 2010, *A&A*, 518, L43
- Greve, T. R., Leonidaki, I., Xilouris, E. M., et al. 2014, *ApJ*, 794, 142
- Greve, T. R., Vieira, J. D., Weiß, A., et al. 2012, *ApJ*, 756, 101
- Gullberg, B., De Breuck, C., Vieira, J. D., et al. 2015, *MNRAS*, 449, 2883
- Hayward, C. C., Jonsson, P., Kereš, D., et al. 2012, *MNRAS*, 424, 951
- Hayward, C. C., Lanz, L., Ashby, M. L. N., et al. 2014, *MNRAS*, 445, 1598
- Helou, G., Soifer, B. T., & Rowan-Robinson, M. 1985, *ApJL*, 298, L7
- Hezaveh, Y. D., Dalal, N., Marrone, D. P., et al. 2016, *ApJ*, 823, 37
- Hezaveh, Y. D., Marrone, D. P., & Holder, G. P. 2012, *ApJ*, 761, 20
- Kennicutt, R. C., & Evans, N. J. 2012, *ARA&A*, 50, 531
- Kennicutt, R. C., Jr. 1998, *ARA&A*, 36, 189
- Kennicutt, R. C., Jr., Hao, C.-N., Calzetti, D., et al. 2009, *ApJ*, 703, 1672
- Koss, M., Mushotzky, R., Baumgartner, W., et al. 2013, *ApJL*, 765, L26
- Lagache, G., Cousin, M., & Chatzikos, M. 2018, *A&A*, 609, A130
- Li, A., & Draine, B. T. 2001, *ApJ*, 554, 778
- Liu, D., Gao, Y., Isaak, K., et al. 2015, *ApJL*, 810, L14
- Liu, L., Weiß, A., Perez-Beaupuits, J. P., et al. 2017, *ApJ*, 846, 5
- Lu, N., Zhao, Y., Díaz-Santos, T., et al. 2017, *ApJS*, 230, 1
- Lu, N., Zhao, Y., Xu, C. K., et al. 2015, *ApJL*, 802, L11
- McMullin, J. P., Waters, B., Schiebel, D., Young, W., & Golap, K. 2007, in ASP Conf. Ser. 376, *Astronomical Data Analysis Software and Systems XVI*, ed. R. A. Shaw, F. Hill, & D. J. Bell (San Francisco, CA: ASP), 127
- Mocanu, L. M., Crawford, T. M., Vieira, J. D., et al. 2013, *ApJ*, 779, 61
- Murphy, E. J., Condon, J. J., Schinnerer, E., et al. 2011, *ApJ*, 737, 67
- Narayanan, D., Cox, T. J., Shirley, Y., et al. 2008, *ApJ*, 684, 996
- Narayanan, D., & Krumholz, M. R. 2014, *MNRAS*, 442, 1411
- Narayanan, D., & Krumholz, M. R. 2017, *MNRAS*, 467, 50
- Narayanan, D., Turk, M., Feldmann, R., et al. 2015, *Natur*, 525, 496
- Neufeld, D. A., Lepp, S., & Melnick, G. J. 1995, *ApJS*, 100, 132
- Omont, A. 2007, *RPPh*, 70, 1099
- Omont, A., Neri, R., Cox, P., et al. 2011, *A&A*, 530, L3
- Omont, A., Yang, C., Cox, P., et al. 2013, *A&A*, 551, A115
- Papadopoulos, P. P., van der Werf, P., Isaak, K., & Xilouris, E. M. 2010, *ApJ*, 715, 775
- Planck Collaboration, Ade, P. A. R., Aghanim, N., et al. 2016, *A&A*, 594, A13
- Rangwala, N., Maloney, P. R., Glenn, J., et al. 2011, *ApJ*, 743, 94
- Riechers, D. A., Bradford, C. M., Clements, D. L., et al. 2013, *Natur*, 496, 329
- Riechers, D. A., Walter, F., Carilli, C. L., & Lewis, G. F. 2009, *ApJ*, 690, 463
- Sanders, D. B., Mazzarella, J. M., Kim, D.-C., Surace, J. A., & Soifer, B. T. 2003, *AJ*, 126, 1607
- Shirley, Y. L. 2015, *PASP*, 127, 299
- Soifer, B. T., Neugebauer, G., & Houck, J. R. 1987, *ARA&A*, 25, 187
- Solomon, P., Vanden Bout, P., Carilli, C., & Guelin, M. 2003, *Natur*, 426, 636
- Solomon, P. M., Downes, D., Radford, S. J. E., & Barrett, J. W. 1997, *ApJ*, 478, 144
- Solomon, P. M., & Vanden Bout, P. A. 2005, *ARA&A*, 43, 677
- Spilker, J. S., Aravena, M., Marrone, D. P., et al. 2015, *ApJ*, 811, 124
- Spilker, J. S., Marrone, D. P., Aguirre, J. E., et al. 2014, *ApJ*, 785, 149
- Spilker, J. S., Marrone, D. P., Aravena, M., et al. 2016, *ApJ*, 826, 112
- Strandet, M. L., Weiss, A., Vieira, J. D., et al. 2016, *ApJ*, 822, 80
- Tadaki, K.-i., Kodama, T., Nelson, E. J., et al. 2017, *ApJL*, 841, L25
- Tielens, A. G. G. M. 2013, *RvMP*, 85, 1021
- van der Werf, P. P., Berciano Alba, A., Spaans, M., et al. 2011, *ApJL*, 741, L38
- van der Werf, P. P., Isaak, K. G., Meijerink, R., et al. 2010, *A&A*, 518, L42
- Vassilev, V., Meledin, D., Lapkin, I., et al. 2008, *A&A*, 490, 1157
- Venturini, S., & Solomon, P. M. 2003, *ApJ*, 590, 740
- Vieira, J. D., Crawford, T. M., Switzer, E. R., et al. 2010, *ApJ*, 719, 763
- Vieira, J. D., Marrone, D. P., Chapman, S. C., et al. 2013, *Natur*, 495, 344
- Weiß, A., De Breuck, C., Marrone, D. P., et al. 2013, *ApJ*, 767, 88
- Weiß, A., Downes, D., Neri, R., et al. 2007, *A&A*, 467, 955
- Weiß, A., Henkel, C., Downes, D., & Walter, F. 2003, *A&A*, 409, L41
- Weiß, A., Requena-Torres, M. A., Güsten, R., et al. 2010, *A&A*, 521, L1
- Yang, C., Gao, Y., Omont, A., et al. 2013, *ApJL*, 771, L24
- Yang, C., Gavazzi, R., Beelen, A., et al. 2019, *A&A*, 624, A138
- Yang, C., Omont, A., Beelen, A., et al. 2016, *A&A*, 595, A80
- Yang, C., Omont, A., Beelen, A., et al. 2017, *A&A*, 608, A144
- Zhang, Z.-Y., Gao, Y., Henkel, C., et al. 2014, *ApJL*, 784, L31